

A *SPITZER* IRS LOW-RESOLUTION SPECTROSCOPIC SEARCH FOR BURIED AGNs IN NEARBY ULTRALUMINOUS INFRARED GALAXIES: A CONSTRAINT ON GEOMETRY BETWEEN ENERGY SOURCES AND DUST

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Received 2006 October 24; accepted 2007 February 5

ABSTRACT

We present the results of *Spitzer* Infrared Spectrograph low-resolution infrared 5–35 μm spectroscopy of nearby ultraluminous infrared galaxies (ULIRGs) at $z < 0.15$. We focus on the search for the signatures of buried active galactic nuclei (AGNs) in the complete sample of ULIRGs classified optically as non-Seyferts (LINERs or H II regions). In addition to polycyclic aromatic hydrocarbon (PAH) emission features at 6.2, 7.7, and 11.3 μm , the conventional tool of starburst-AGN separation, we use the optical depths of the 9.7 and 18 μm silicate dust absorption features to infer the geometry of energy sources and dust at the nuclei of these ULIRGs, namely, whether the energy sources are spatially well mixed with dust (a normal starburst) or are more centrally concentrated than the dust (a buried AGN). Infrared spectra of at least 30%, and possibly 50%, of the observed optical non-Seyfert ULIRGs are naturally explained by emission consisting of (1) energetically insignificant, modestly obscured ($A_V < 20$ –30 mag) PAH-emitting normal starbursts and (2) energetically dominant, highly dust-obscured, centrally concentrated energy sources with no PAH emission. We interpret the latter component as a buried AGN. The fraction of ULIRGs showing some buried AGN signatures is higher in LINER ULIRGs than in H II region ULIRGs. Most of the luminous buried AGN candidates are found in ULIRGs with cool far-infrared colors. Where the absorption-corrected intrinsic AGN luminosities are derivable with little uncertainty, they are found to be of the order of $10^{12} L_\odot$, accounting for the bulk of the ULIRGs' luminosities. The 5–35 μm spectroscopic starburst/AGN classifications are generally consistent with our previous classifications based on 3–4 μm spectra for the same sample.

Subject headings: galaxies: active — galaxies: ISM — galaxies: nuclei — galaxies: Seyfert — galaxies: starburst — infrared: galaxies

1. INTRODUCTION

Galaxies that radiate very large luminosities as infrared dust emission ($L_{\text{IR}} \gtrsim 10^{12} L_\odot$) are called ultraluminous infrared galaxies (ULIRGs; Sanders & Mirabel 1996). The huge infrared luminosities mean that (1) very powerful energy sources, starbursts and/or active galactic nuclei (AGNs), are present hidden behind dust; (2) energetic radiation from the energy sources is absorbed by the surrounding dust; and (3) the heated dust grains re-emit this energy as infrared thermal radiation. In the local universe ($z < 0.3$), the ULIRG population dominates the bright end of the luminosity function (Soifer et al. 1987). The impor-

tance of the ULIRG population i.e., the comoving infrared energy density, increases rapidly with increasing redshift (Le Floc'h et al. 2005). Identifying the dust-obscured energy sources of nearby ULIRGs is thus very important, not only to unveil the nature of nearby ULIRGs, but also to clarify the obscured AGN-starburst connections in the high-redshift ($z > 0.5$) ULIRG population, where detailed investigations are difficult with existing facilities.

In nearby ULIRGs, it is now widely accepted that the dominant energy sources are compact (< 500 pc), highly dust-obscured nuclear cores, rather than extended ($> \text{kpc}$), weakly obscured starburst activity in the host galaxies (Soifer et al. 2000; Fischer 2000). The most important issue is to determine whether the obscured compact cores are powered by very compact starbursts and/or AGNs, from which strong emission is produced in very compact regions around the central accreting supermassive black holes.

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Report Documentation Page				Form Approved OMB No. 0704-0188	
Public reporting burden for the collection of information is estimated to average 1 hour per response, including the time for reviewing instructions, searching existing data sources, gathering and maintaining the data needed, and completing and reviewing the collection of information. Send comments regarding this burden estimate or any other aspect of this collection of information, including suggestions for reducing this burden, to Washington Headquarters Services, Directorate for Information Operations and Reports, 1215 Jefferson Davis Highway, Suite 1204, Arlington VA 22202-4302. Respondents should be aware that notwithstanding any other provision of law, no person shall be subject to a penalty for failing to comply with a collection of information if it does not display a currently valid OMB control number.					
1. REPORT DATE FEB 2007		2. REPORT TYPE		3. DATES COVERED 00-00-2007 to 00-00-2007	
4. TITLE AND SUBTITLE A Spitzer IRS Low-Resolution Spectroscopic Search for Buried AGNs in Nearby Ultraluminous Infrared Galaxies: A Constraint on Geometry Between Energy Sources and Dust				5a. CONTRACT NUMBER	
				5b. GRANT NUMBER	
				5c. PROGRAM ELEMENT NUMBER	
6. AUTHOR(S)				5d. PROJECT NUMBER	
				5e. TASK NUMBER	
				5f. WORK UNIT NUMBER	
7. PERFORMING ORGANIZATION NAME(S) AND ADDRESS(ES) Naval Research Laboratory, Remote Sensing Division, Code 7211, 4555 Overlook Avenue SW, Washington, DC, 20375				8. PERFORMING ORGANIZATION REPORT NUMBER	
9. SPONSORING/MONITORING AGENCY NAME(S) AND ADDRESS(ES)				10. SPONSOR/MONITOR'S ACRONYM(S)	
				11. SPONSOR/MONITOR'S REPORT NUMBER(S)	
12. DISTRIBUTION/AVAILABILITY STATEMENT Approved for public release; distribution unlimited					
13. SUPPLEMENTARY NOTES					
14. ABSTRACT see report					
15. SUBJECT TERMS					
16. SECURITY CLASSIFICATION OF:			17. LIMITATION OF ABSTRACT Same as Report (SAR)	18. NUMBER OF PAGES 29	19a. NAME OF RESPONSIBLE PERSON
a. REPORT unclassified	b. ABSTRACT unclassified	c. THIS PAGE unclassified			

If a luminous AGN is present hidden behind dust in a *torus* geometry, as inferred for the classical Seyfert population (e.g., Antonucci 1993), a large amount of energetic AGN radiation can escape along the direction perpendicular to the torus and form the so-called narrow-line regions (NLRs; Antonucci 1993). NLRs are the sources of strong forbidden emission lines seen in the optical–infrared, and their relative strengths are different from those of a starburst because the spectrum at the ionizing high-energy UV range ($\lambda < 912 \text{ \AA}$) is harder in an AGN than in a starburst. An AGN obscured by a torus-shaped dust distribution, with a *well-developed NLR*, is thus distinguishable from a starburst galaxy relatively easily, through optical spectroscopy (classified as Seyfert 2; Veilleux & Osterbrock 1987) or high-resolution infrared spectroscopy (Genzel et al. 1998; Armus et al. 2007).

Since the gas and dust in an AGN have angular momentum with respect to the central supermassive black hole, an axisymmetric spatial distribution is more natural than spherical geometry. In this case, the column densities can be high in certain directions (the torus directions) but low in others (along the torus axis). The classical Seyfert-type AGN is naturally understood as the population where most of the torus axis direction is transparent to the AGN’s ionizing UV–soft-X-ray radiation. However, compared to the classical Seyfert AGNs, ULIRGs are known to contain a substantially larger amount of gas and dust in their nuclei (Sanders & Mirabel 1996). For a fixed angular momentum, the increased amount of gas and dust results in higher column densities in all directions. If high-density gas/dust blocks the bulk of the AGN’s ionizing radiation at very small radii ($\ll 10 \text{ pc}$) even in the torus axis direction, then no significant NLR will be present. It is very likely that any luminous AGNs in many ULIRGs exist in this *buried* condition, with no obvious Seyfert signatures in the optical spectra (hereafter “optical non-Seyfert”). Studying only AGNs with well-developed NLRs provides a very incomplete picture of the AGN population and will miss many *buried* AGNs, whose understanding may be essential to clarify the true nature of ULIRGs.

An often-used indicator to separate a buried AGN from a normal starburst is the emission from polycyclic aromatic hydrocarbons (PAHs), seen in infrared spectra at $\lambda_{\text{rest}} = 3\text{--}25 \text{ }\mu\text{m}$ in the rest frame (Genzel et al. 1998; Imanishi & Dudley 2000). In a normal starburst, UV-emitting H II regions and molecular gas and dust are spatially well mixed. PAHs are excited by far-UV photons from stars, and strong PAH emission is produced in photo-dissociation regions (PDRs), the interfaces between H II regions and molecular gas (Sellgren 1981), unless the metallicity is very low (Houck et al. 2004b). In a normal starburst, the equivalent width of the PAH emission is insensitive to level of dust extinction, as both the PAH emission and the nearby continuum are similarly attenuated. Thus, a normal starburst should always show large equivalent width PAH emission, regardless of dust extinction. Near an AGN, the PAHs are destroyed by strong X-ray radiation from the AGN (Voit 1992; Siebenmorgen et al. 2004). If the PAHs are sufficiently shielded by obscuration from the X-ray emission of the central AGN, the PAHs can survive. However, the PAH-exciting far-UV emission from the AGN will also be attenuated by the obscuration, so in a pure AGN without starburst activity, the PAH emission will be very weak. Instead, a PAH-free continuum from hot, submicron-sized dust grains heated by the AGN is observed. In a starburst/AGN composite galaxy, PAH emission from the starburst regions is seen, but its equivalent width will be smaller than in a pure starburst, because of the dilution by PAH-free continuum produced by the AGN. Thus, this PAH-based method can be used to disentangle a buried AGN from a normal starburst at an obscured ULIRG’s core. Since the PAH emission features are spectrally very broad, low-resolution in-

frared spectroscopy is adequate. In addition, ULIRGs are by definition bright in the infrared, so that this infrared PAH technique can be applied to a large number of ULIRGs in a reasonable amount of telescope time, and thus can be a very effective tool to systematically identify the hidden energy sources of ULIRGs’ cores.

Using low-resolution infrared spectroscopy, we also have another powerful tool to distinguish buried AGNs from normal starbursts in the cores of ULIRGs, by investigating the *geometry* between energy sources and the gas and dust. In a normal starburst, the energy sources (stars) and gas/dust are spatially well mixed (Fig. 1a) (Puxley 1991; McLeod et al. 1993; Forster Schreiber et al. 2001), while in a buried AGN, the energy source (the central accreting supermassive black hole) is very compact and is more centrally concentrated than the surrounding gas and dust (Fig. 1b) (Soifer et al. 2000; Imanishi & Maloney 2003; Siebenmorgen et al. 2004). The difference in geometry is reflected in two different ways in the observed low-resolution infrared spectra. The first is that, while the absolute optical depths of dust absorption features in the $3\text{--}10 \text{ }\mu\text{m}$ range cannot exceed a certain threshold in a normal starburst with mixed dust/source geometry (Fig. 1a), they can be arbitrarily large in a buried AGN (Fig. 1b) (Imanishi & Maloney 2003; Imanishi et al. 2006a; Levenson et al. 2007). The second is that a strong dust temperature gradient is found in a buried AGN (Fig. 2), but not in a normal starburst. This dust temperature gradient can be investigated by comparing optical depths of dust absorption features at different infrared wavelengths (Dudley & Wynn-Williams 1997; Imanishi 2000; Imanishi et al. 2006a).

In this paper we present low-resolution $5\text{--}35 \text{ }\mu\text{m}$ spectra of nearby ULIRGs at $z < 0.15$ to search for the signatures of optically elusive (Maiolino et al. 2003) buried AGNs in the cores of ULIRGs and to determine the energetic contribution of buried AGNs. Throughout this paper, $H_0 = 75 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$, and $\Omega_\Lambda = 0.7$ are adopted. The physical scale of $1''$ is 0.34 kpc in the nearest source at $z = 0.018$, 1.72 kpc at $z = 0.1$, and 2.44 kpc in the furthest source at $z = 0.15$.

2. TARGETS

Our main targets are optically non-Seyfert ULIRGs at $z < 0.15$ in the *IRAS* 1 Jy sample (Kim & Sanders 1998). In this 1 Jy sample, there are 69 ULIRGs at $z < 0.15$. Based on the optical classifications by Veilleux et al. (1999a, their Table 2), 28 ULIRGs are classified optically as LINERs and 20 ULIRGs are classified optically as H II regions. These 48 ($=28 + 20$) ULIRGs are selected, comprising $\sim 70\%$ of the 1 Jy sample at $z < 0.15$. The remaining $\sim 30\%$ ULIRGs are classified optically as Seyferts (17 sources) and optically unclassified (4 sources; Veilleux et al. 1999a). Table 1 summarizes the basic information and *IRAS*-based infrared properties of the selected LINER and H II region ULIRGs. Low-resolution infrared $5\text{--}35 \text{ }\mu\text{m}$ spectra for this complete sample are presented in this paper. For most sources, ancillary ground-based infrared $3\text{--}4 \text{ }\mu\text{m}$ spectra are available (Imanishi et al. 2006a).

3. OBSERVATIONS AND DATA ANALYSIS

Observations of all the 48 ULIRGs were performed using the Infrared Spectrograph (IRS; Houck et al. 2004a) on board *Spitzer* (Werner et al. 2004). All four modules, Short-Low 2 (SL2; $5.2\text{--}7.7 \text{ }\mu\text{m}$) and 1 (SL1; $7.4\text{--}14.5 \text{ }\mu\text{m}$), Long-Low 2 (LL2; $14.0\text{--}21.3 \text{ }\mu\text{m}$) and 1 (LL1; $19.5\text{--}38.0 \text{ }\mu\text{m}$) were used, to obtain full $5\text{--}35 \text{ }\mu\text{m}$ low-resolution ($R \sim 100$) spectra. 13 out of the 48 sources were in the GTO target list (PI: J. Houck, pid = 105), and their archival spectra were analyzed. The remaining 35 sources were observed through our GO-1 program

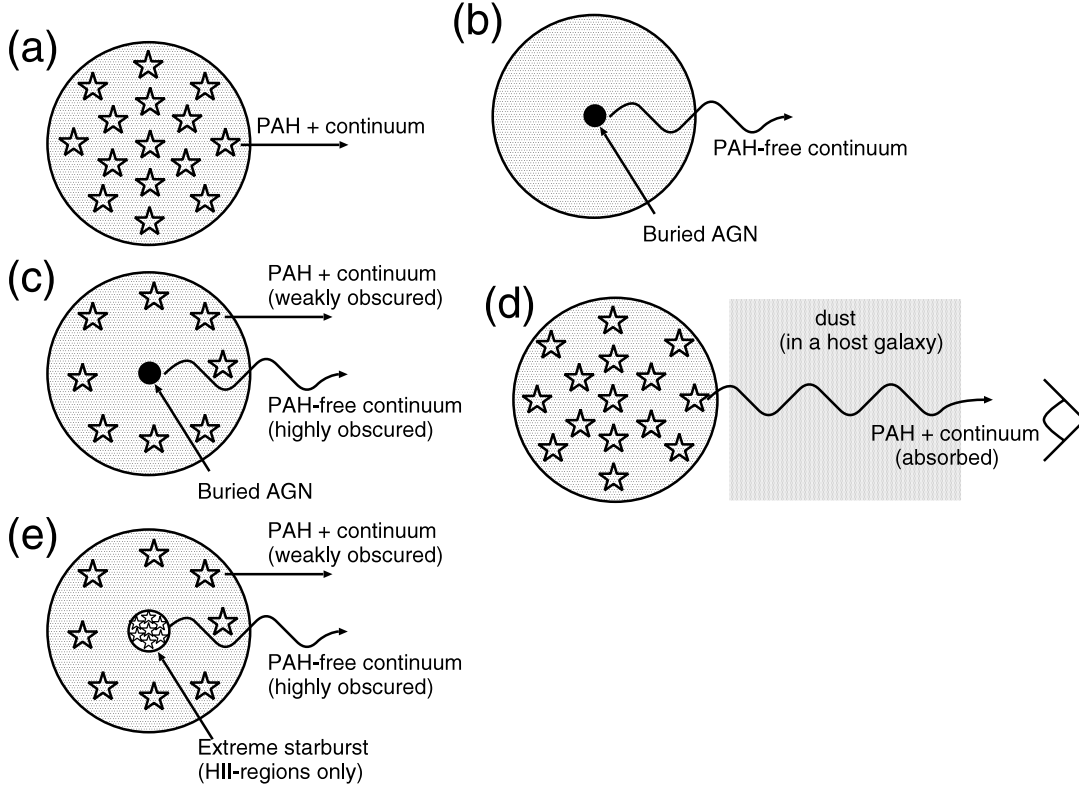


FIG. 1.—(a) Geometry of energy sources and dust in a normal starburst. The open star symbols indicate “stars” in a starburst. The energy sources (*stars*) and dust are spatially well mixed. Large equivalent width PAH emission is observed. (b) Geometry of the energy source and dust in a buried AGN. The energy source (a mass-accreting supermassive black hole) is more centrally concentrated than the surrounding dust. No PAH emission is found. (c) A buried AGN/starburst composite. The starburst surrounds the central buried AGN. The observed spectrum is a superposition of PAH + continuum emission from the starburst and PAH-free continuum from the AGN. Since the buried AGN is more highly obscured than the surrounding starbursts, the starburst emission generally makes a strong contribution to an observed flux even if the starburst is energetically insignificant. (d) A normal starburst nucleus (mixed dust/source geometry) obscured by a large amount of foreground dust in a host galaxy. This geometry can produce a strong dust absorption feature whose optical depth is larger than the maximum threshold obtained in the mixed dust/source geometry. (e) An exceptionally centrally concentrated extreme starburst, whose emitting volume is predominantly occupied by H II regions. Such an extreme starburst produces no PAH emission, like a buried AGN.

(PI: M. Imanishi, pid = 2306) and another GO-1 program (PI: S. Veilleux, pid = 3187). The GO-1 program by Dr. Veilleux covered only SL2 and SL1. For these ULIRGs, LL2 and LL1 spectra were taken in our GO-1 program.

For our program, each observation consists of $60 \text{ s} \times 2$ cycles for SL and $30 \text{ s} \times 4$ cycles for LL. The observation was performed both at A and B beam positions, and so total exposure time was 240 s for both SL and LL. For the programs by Houck and Veilleux, the total exposure times are similar or smaller for bright sources. Table 2 tabulates the observing log of the 48 ULIRGs. The slit width is $3.6''$ or 2 pixels for SL2 ($1.8'' \text{ pixel}^{-1}$) and is $3.7''$ or ~ 2 pixels for SL1 ($1.8'' \text{ pixel}^{-1}$). For LL2 and LL1, the slit widths are $10.5''$ and $10.7''$, respectively, which correspond to ~ 2 pixels for both LL2 ($5.1'' \text{ pixel}^{-1}$) and LL1 ($5.1'' \text{ pixel}^{-1}$).

The latest pipeline-processed data products at the time of our data analysis (S11–14, pbcd files) were used. Frames taken at the A position were subtracted from those taken at the B position, to remove background emission, mostly the zodiacal light. Then, spectra were extracted in a standard manner using the task `apa11` in IRAF.³ Apertures with 4–6 and 4–5 pixels were employed for SL and LL data, respectively, depending on the spatial extent of individual sources. Then, spectra extracted for the A and B positions were summed. Wavelength calibration was made based on the files of

the *Spitzer* pipeline processed data, named “b0_wavsamp.tbl” and “b2_wavsamp.tbl” for SL and LL, respectively. It is believed to be accurate within $0.1 \mu\text{m}$. A small level of possible error in wavelength calibration will not affect our main discussions. Since emission from all ULIRGs is dominated by spatially compact sources at the observed wavelength, flux calibration was performed using the *Spitzer* pipeline processed files “b0_fluxcon.tbl” (SL) and “b2_fluxcon.tbl” (LL). For SL1 spectra, data at $\lambda_{\text{obs}} > 14.5 \mu\text{m}$ in the observed frame are invalid (Infrared Spectrograph Data Handbook ver. 1.0) and so were discarded. For LL1 spectra, we used only data at $\lambda_{\text{obs}} < 35 \mu\text{m}$, because the scatter is large at $\lambda_{\text{obs}} > 35 \mu\text{m}$. No defringing was attempted.

For flux calibration, we adopted the values of the pipeline processed data. We made no attempt to recalibrate our spectra using *IRAS* measurements at 12 and $25 \mu\text{m}$, primarily because (1) only upper limits are provided for *IRAS* $12 \mu\text{m}$ photometry in a sizable fraction ($>60\%$; 31/48) of the observed ULIRGs and for $25 \mu\text{m}$ data in some fraction ($\sim 20\%$; 10/48) of them (Table 1), and (2) the validity of recalibrating IRS fixed-aperture slit spectra, using the large aperture photometry of *IRAS*, is not obvious for ULIRGs that possibly show a non-negligible fraction of spatially extended dim emission outside the IRS slit. Hence, the absolute flux calibration is dependent on the accuracy of the pipeline processed data, which is taken to be $<20\%$ for SL and LL (Infrared Spectrograph Data Handbook). This level of possible flux uncertainty will not affect our main discussions significantly. In fact, in the majority of ULIRGs with *IRAS* detection at $25 \mu\text{m}$,

³ IRAF is distributed by the National Optical Astronomy Observatories, which are operated by the Association of Universities for Research in Astronomy, Inc. (AURA), under cooperative agreement with the National Science Foundation.

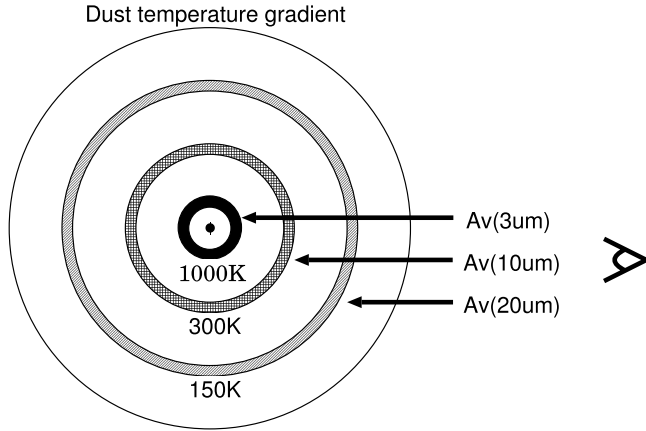


FIG. 2.—Schematic diagram of a strong dust temperature gradient around a centrally concentrated energy source (*central small filled circle*). Inner dust close to the central energy source has higher temperature than dust far away from the energy source. Assuming approximately blackbody radiation, dust at 1000 K dust, close to the innermost dust sublimation radius, shows an emission peak at $\lambda_{\text{rest}} \sim 3 \mu\text{m}$. More distant 300 K dust and even 150 K dust have emission peaks at $\lambda_{\text{rest}} \sim 10$ and $\sim 20 \mu\text{m}$, respectively. This is a very simplified picture, and in an actual case, dust with other temperatures can also contribute to the observed flux at each wavelength. For any reasonable dust radial density distribution, it is always true that the contribution from inner, hotter dust emission is larger at a shorter wavelength in this wavelength range. When we look at the energy source from the right side, an observed flux is dominated by dust at the foreground side (*right*), rather than dust with the same temperature at the distant side (*left*), due to flux attenuation, particularly for outer cooler dust. Thus, the relation $A_V(3 \mu\text{m}) > A_V(10 \mu\text{m}) > A_V(20 \mu\text{m})$ should hold. Dust extinction toward the $3 \mu\text{m}$ continuum-emitting region, $A_V(3 \mu\text{m})$, can be estimated from the optical depths of absorption features at $3.4 \mu\text{m}$ due to bare carbonaceous dust and at $3.1 \mu\text{m}$ due to ice-covered dust (Imanishi et al. 2001, 2006a; Imanishi 2006). The extinction toward the 10 and $20 \mu\text{m}$ continuum-emitting regions, $A_V(10 \mu\text{m})$ and $A_V(20 \mu\text{m})$, can be estimated from the optical depths of 9.7 and $18 \mu\text{m}$ silicate dust absorption features in *Spitzer* IRS spectra, respectively. Although a spherical geometry is illustrated here for simplicity, the essence is unchanged as long as a centrally concentrated compact energy source is obscured by dust along our line of sight.

the *Spitzer* IRS $25 \mu\text{m}$ flux agrees within 30% to the *IRAS* $25 \mu\text{m}$ data, and the *Spitzer* IRS flux is generally smaller. In a small fraction of ULIRGs, *Spitzer* IRS $25 \mu\text{m}$ flux is $\sim 50\%$ of *IRAS* $25 \mu\text{m}$ data, possibly because extended dim emission is important. For ULIRGs with *IRAS* nondetection at $25 \mu\text{m}$, the measured *Spitzer* IRS flux at $25 \mu\text{m}$ is always smaller than the upper limit of the *IRAS* data.

For a fraction of the ULIRGs, slight flux discrepancies between SL1 and LL2 were discernible, ranging from 10% to 40%. When the discrepancy is seen, the SL1 flux ($3.7''$ wide slit) is usually smaller than the LL2 flux ($10.5''$), suggesting the presence of extended emission in these ULIRGs. In these cases we adjusted the smaller SL1 (and SL2) flux to match the larger LL2 flux.

The main part of the $9.7 \mu\text{m}$ silicate absorption feature extends $\lambda_{\text{rest}} = 8\text{--}13 \mu\text{m}$ in the rest wavelength (e.g., Chiar & Tielens 2006). For ULIRGs at $z > 0.11$, the feature is not fully covered with SL1 spectra at $\lambda_{\text{obs}} = 7.4\text{--}14.5 \mu\text{m}$. Even for ULIRGs at $z < 0.11$, the continuum data at $\lambda_{\text{rest}} = 13\text{--}14 \mu\text{m}$, which are required to estimate the strength of the $9.7 \mu\text{m}$ silicate absorption, fall beyond $\lambda_{\text{obs}} > 14.5 \mu\text{m}$, outside the wavelength range of SL1. Thus, for the majority of the observed ULIRGs, LL2 spectra are essential to measure the optical depth of the $9.7 \mu\text{m}$ silicate absorption feature with a small ambiguity.

Appropriate spectral binning with 2 or 4 pixels was applied to reduce the scatter of data points at SL2 ($5.2\text{--}7.7 \mu\text{m}$) for some faint ULIRGs, and at $\lambda_{\text{obs}} \sim 10 \mu\text{m}$ for ULIRGs displaying very strong $9.7 \mu\text{m}$ silicate absorption features.

4. RESULT

Figures 3 and 4 present low-resolution $5\text{--}35 \mu\text{m}$ spectra of ULIRGs classified optically as LINERs and H II regions, respectively. For most of the sources, full $5\text{--}35 \mu\text{m}$ spectra are shown here for the first time. Although spectra of the five ULIRGs in the GTO program (Arp 220, IRAS 08572+3915, 12112+0305, 14348–1447, and 22491–1808) were presented elsewhere (Spoon et al. 2006; Armus et al. 2007), we include their spectra to comprise a complete sample, to which our detailed and uniform analysis is applied. For these ULIRGs, their and our spectra are, in overall, consistent with each other.

The spectra in Figures 3 and 4 are suitable for displaying the properties of the 9.7 and $18 \mu\text{m}$ silicate dust absorption features. However, they are not very useful for PAH emission features. Figure 5 presents zoom-in spectra at $\lambda_{\text{obs}} = 5.2\text{--}14.5 \mu\text{m}$ of optical LINER and H II region ULIRGs to better examine the properties of the PAH emission features.

4.1. PAH Emission

The majority of ULIRGs in Figure 5 show clearly detectable PAH emission features at $\lambda_{\text{rest}} = 6.2, 7.7$, and $11.3 \mu\text{m}$. A fraction of the ULIRGs also display the PAH emission feature at $\lambda_{\text{rest}} = 8.6 \mu\text{m}$. Emission at $\lambda_{\text{rest}} = 12.8 \mu\text{m}$ is also detected (Figs. 3 and 4), but it can be attributed to both the $12.8 \mu\text{m}$ PAH and [Ne II] emission lines, which are not clearly resolved in our low-resolution spectra. We thus will not discuss the $12.8 \mu\text{m}$ PAH emission feature further.

The $6.2 \mu\text{m}$ PAH emission feature is attributed to an aromatic C-C stretching fundamental (Allamandola et al. 1989). It is relatively weak but is isolated both from other PAH emission and from the strong $9.7 \mu\text{m}$ silicate dust absorption feature. Although absorption at $6.0 \mu\text{m}$ due to H_2O ice and at $6.85 \mu\text{m}$ by hydrogenated amorphous carbon (HAC) are present (Spoon et al. 2002), their relatively weak nature, compared to the $9.7 \mu\text{m}$ silicate absorption, makes it possible to measure the $6.2 \mu\text{m}$ PAH emission strength with little ambiguity. For this reason, the $6.2 \mu\text{m}$ feature has often been used to measure the PAH emission strength in a ULIRG (Fischer 2000; Spoon et al. 2007). To estimate the strength of the $6.2 \mu\text{m}$ PAH emission feature, we adopt a linear continuum determined from data at $\lambda_{\text{rest}} = 6.1$ and $6.45 \mu\text{m}$, which is shown as solid lines in Figure 5.

The $11.3 \mu\text{m}$ PAH emission is due to an out-of-plane aromatic C-H bending vibration (Allamandola et al. 1989). Since it is situated close to the peak of the strong $9.7 \mu\text{m}$ silicate absorption, its flux attenuation can be severe if the PAH-emitting regions are highly dust obscured. Thus, its flux relative to other PAH features can be a good indicator of the dust obscuration of the PAH-emitting regions. To estimate the $11.3 \mu\text{m}$ PAH strength, we adopt a linear continuum connecting data points at $\lambda_{\text{rest}} = 11.0$ and $11.6 \mu\text{m}$, which is shown as solid lines in Figure 5.

The $7.7 \mu\text{m}$ PAH feature is due to aromatic C-C stretching vibrations (Allamandola et al. 1989). Although it is the strongest PAH emission feature, its substantial wavelength overlap with the broad, strong $9.7 \mu\text{m}$ silicate dust absorption feature makes it very difficult to distinguish between $7.7 \mu\text{m}$ PAH emission and $9.7 \mu\text{m}$ silicate absorption, particularly for sources showing strong silicate absorption, such as the majority of ULIRGs. To reduce the effects of the $9.7 \mu\text{m}$ absorption feature, we adopt an excess above a linear continuum determined from data at $\lambda_{\text{rest}} = 7.3$ and $8.1 \mu\text{m}$ (shown as solid lines in Fig. 5) as the $7.7 \mu\text{m}$ PAH emission. Since many previous papers (e.g., Genzel et al. 1998) measured the $7.7 \mu\text{m}$ PAH fluxes based on different definitions, readers must

TABLE 1
OBSERVED ULIRGS AT $z < 0.15$ AND THEIR *IRAS*-BASED INFRARED EMISSION PROPERTIES

Object (1)	Redshift (2)	f_{12} (Jy) (3)	f_{25} (Jy) (4)	f_{60} (Jy) (5)	f_{100} (Jy) (6)	$\log L_{\text{IR}}$ ($L_{\odot}$) (7)	f_{25}/f_{60} (8)	Optical Class (9)
IRAS 00188+0856	0.128	<0.12	0.37	2.59	3.40	12.3	0.14 (C)	LINER
IRAS 00482+2721	0.129	<0.10	<0.18	1.13	1.84	12.0	<0.16 (C)	LINER
IRAS 03250+1606	0.129	<0.10	<0.15	1.38	1.77	12.1	<0.11 (C)	LINER
IRAS 04103+2838	0.118	0.08	0.54	1.82	1.71	12.2	0.30 (W)	LINER
IRAS 08572+3915	0.058	0.32	1.70	7.43	4.59	12.1	0.23 (W)	LINER
IRAS 09039+0503	0.125	0.07	0.12	1.48	2.06	12.1	0.08 (C)	LINER
IRAS 09116+0334	0.146	<0.09	<0.14	1.09	1.82	12.1	<0.13 (C)	LINER
IRAS 09539+0857	0.129	<0.15	<0.15	1.44	1.04	12.0	<0.11 (C)	LINER
IRAS 10378+1108	0.136	<0.11	0.24	2.28	1.82	12.3	0.11 (C)	LINER
IRAS 10485+1447	0.133	<0.11	0.25	1.73	1.66	12.2	0.14 (C)	LINER
IRAS 10494+4424	0.092	<0.12	0.16	3.53	5.41	12.1	0.05 (C)	LINER
IRAS 11095+0238	0.106	0.06	0.42	3.25	2.53	12.2	0.13 (C)	LINER
IRAS 11130+2659	0.136	<0.09	0.20	1.21	1.24	12.1	0.17 (C)	LINER
IRAS 12112+0305	0.073	0.12	0.51	8.50	9.98	12.3	0.06 (C)	LINER
IRAS 12127+1412	0.133	<0.13	0.24	1.54	1.13	12.1	0.16 (C)	LINER
IRAS 12359+0725	0.138	0.09	0.15	1.33	1.12	12.1	0.11 (C)	LINER
IRAS 13335+2612	0.125	<0.13	<0.14	1.40	2.10	12.1	<0.10 (C)	LINER
IRAS 14252+1550	0.149	<0.09	<0.23	1.15	1.86	12.2	<0.20 (C)	LINER
IRAS 14348+1447	0.083	0.07	0.49	6.87	7.07	12.3	0.07 (C)	LINER
IRAS 15327+2340 (Arp 220)	0.018	0.48	7.92	103.33	112.40	12.1	0.08 (C)	LINER
IRAS 16090+0139	0.134	0.09	0.26	3.61	4.87	12.5	0.07 (C)	LINER
IRAS 16468+5200	0.150	<0.06	0.10	1.01	1.04	12.1	0.10 (C)	LINER
IRAS 16487+5447	0.104	<0.07	0.20	2.88	3.07	12.1	0.07 (C)	LINER
IRAS 17028+5817	0.106	<0.06	0.10	2.43	3.91	12.1	0.04 (C)	LINER
IRAS 17044+6720	0.135	<0.07	0.36	1.28	0.98	12.1	0.28 (W)	LINER
IRAS 21329+2346	0.125	0.05	0.12	1.65	2.22	12.1	0.07 (C)	LINER
IRAS 23234+0946	0.128	<0.06	0.08	1.56	2.11	12.1	0.05 (C)	LINER
IRAS 23327+2913	0.107	<0.06	0.22	2.10	2.81	12.1	0.10 (C)	LINER
IRAS 00091+0738	0.118	<0.07	0.22	2.63	2.52	12.2	0.08 (C)	H II
IRAS 00456+2904	0.110	<0.08	0.14	2.60	3.38	12.2	0.05 (C)	H II
IRAS 01004+2237	0.118	0.11	0.66	2.29	1.79	12.3	0.29 (W)	H II
IRAS 01166+0844	0.118	0.07	0.17	1.74	1.42	12.1	0.10 (C)	H II
IRAS 01298+0744	0.136	<0.12	0.19	2.47	2.08	12.3	0.08 (C)	H II
IRAS 01569+2939	0.141	<0.11	0.14	1.73	1.51	12.2	0.08 (C)	H II
IRAS 02411+0353	0.144	<0.08	0.22	1.37	1.95	12.2	0.16 (C)	H II
IRAS 10190+1322	0.077	<0.07	0.38	3.33	5.57	12.0	0.11 (C)	H II
IRAS 11387+4116	0.149	0.12	<0.14	1.02	1.51	12.2	<0.14 (C)	H II
IRAS 11506+1331	0.127	<0.10	0.19	2.58	3.32	12.3	0.07 (C)	H II
IRAS 13509+0442	0.136	0.10	<0.23	1.56	2.53	12.2	<0.15 (C)	H II
IRAS 13539+2920	0.108	<0.09	0.12	1.83	2.73	12.0	0.07 (C)	H II
IRAS 14060+2919	0.117	<0.10	0.14	1.61	2.42	12.1	0.09 (C)	H II
IRAS 15206+3342	0.125	0.08	0.35	1.77	1.89	12.2	0.20 (C)	H II
IRAS 15225+2350	0.139	<0.07	0.18	1.30	1.48	12.1	0.14 (C)	H II
IRAS 16474+3430	0.111	<0.13	0.20	2.27	2.88	12.1	0.09 (C)	H II
IRAS 20414+1651	0.086	<0.65	0.35	4.36	5.25	12.2	0.08 (C)	H II
IRAS 21208+0519	0.130	<0.09	<0.15	1.17	1.66	12.0	<0.13 (C)	H II
IRAS 22206+2715	0.132	<0.10	<0.16	1.75	2.33	12.1	<0.09 (C)	H II
IRAS 22491+1808	0.076	0.05	0.55	5.44	4.45	12.1	0.10 (C)	H II

NOTES.—Col. (1): Object name. Col. (2): Redshift. Cols. (3)–(6): f_{12} , f_{25} , f_{60} , and f_{100} are *IRAS* fluxes at 12, 25, 60, and 100 μm , respectively, taken from Kim & Sanders (1998). Col. (7): Decimal logarithm of infrared (8–1000 μm) luminosity in units of solar luminosity ($L_{\odot}$), calculated with $L_{\text{IR}} = 2.1 \times 10^{39} \times D(\text{Mpc})^2 \times (13.48 \times f_{12} + 5.16 \times f_{25} + 2.58 \times f_{60} + f_{100}) \text{ ergs s}^{-1}$ (Sanders & Mirabel 1996). Since the calculation is based on our adopted cosmology, the infrared luminosities slightly (<10%) differ from the values shown in Kim & Sanders (1998, their Table 1, col. [15]). For sources that have upper limits in some *IRAS* bands, we can derive upper and lower limits of the infrared luminosity by assuming that the actual flux is the *IRAS* upper limit and zero value, respectively. The difference of the upper and lower values is usually very small, less than 0.2 dex. We assume that the infrared luminosity is the average of these values. Col. (8): *IRAS* 25–60 μm flux ratio. ULIRGs with $f_{25}/f_{60} < 0.2$ and > 0.2 are classified as cool and warm sources (denoted as “C” and “W”), respectively (Sanders et al. 1988b). Col. (9): Optical spectral classification by Veilleux et al. (1999a).

be careful when our value is compared with other estimates in the literature. In this paper, although the 7.7 μm PAH strength is shown for reference, it will not be used when possible uncertainty coming from the effects of the strong 9.7 μm silicate feature may complicate quantitative discussions.

The 8.6 μm PAH emission due to in-plane aromatic C-H bending vibration (Allamandola et al. 1989) is not used in this paper, because it is weak and overlaps with the 9.7 μm silicate absorption, which make its strength highly uncertain. Other weak PAH emission features may be present (e.g., $\lambda_{\text{rest}} = 5.7 \mu\text{m}$; Spoon

TABLE 2
Spitzer IRS OBSERVING LOG

OBJECT (1)	PID (2)	DATE (UT) (3)	INTEGRATION TIME (s)			
			SL2 (4)	SL1 (5)	LL2 (6)	LL1 (7)
IRAS 00188–0856	105	2003 Dec 17	240	240	180	180
IRAS 00482–2721	3187 + 2306	2005 Jul 7 + 2004 Jul 18	240	240	240	240
IRAS 03250+1606	3187 + 2306	2005 Feb 11	240	240	240	240
IRAS 04103–2838	3187 + 2306	2005 Feb 10 + 2004 Aug 10	240	240	240	240
IRAS 08572+3915	105	2004 Apr 15	84	84	84	84
IRAS 09039+0503	3187 + 2306	2005 Apr 18 + 2004 Nov 17	240	240	240	240
IRAS 09116+0334	2306	2005 Apr 21	240	240	240	240
IRAS 09539+0857	3187 + 2306	2005 Jun 5 + 2005 May 31	240	240	240	240
IRAS 10378+1108	105	2005 Jun 8	240	240	180	180
IRAS 10485–1447	3187 + 2306	2005 May 23 + 2005 Jan 4	240	240	480 ^a	480 ^a
IRAS 10494+4424	2306	2004 Nov 17	240	240	240	240
IRAS 11095–0238	105	2005 Jun 7	240	120	240	120
IRAS 11130–2659	2306	2005 Jul 12	240	240	240	240
IRAS 12112+0305	105	2004 Jan 4	84	84	120	120
IRAS 12127–1412	3187 + 2306	2005 Jun 30	240	240	240	240
IRAS 12359–0725	2306	2005 Jun 30	240	240	240	240
IRAS 13335–2612	3187 + 2306	2005 Feb 15 + 2004 Jul 17	240	240	240	240
IRAS 14252–1550	2306	2004 Jul 17	240	240	240	240
IRAS 14348–1447	105	2004 Feb 7	120	120	120	120
Arp 220	105	2004 Feb 29	84	84	60	60
IRAS 16090–0139	105	2005 Aug 6	120	120	180	180
IRAS 16468+5200	2306	2004 Jul 14	480 ^b	480 ^b	480 ^a	480 ^a
IRAS 16487+5447	2306	2004 Jul 17	240	240	240	240
IRAS 17028+5817	2306	2004 Jul 17	240	240	240	240
IRAS 17044+6720	2306	2004 Jul 17	240	240	240	240
IRAS 21329–2346	3187 + 2306	2004 Nov 16 + 2004 Oct 23	240	240	240	240
IRAS 23234+0946	3187 + 2306	2004 Dec 13 + 2004 Dec 8	240	240	240	240
IRAS 23327+2913	2306	2004 Dec 8	240	240	240	240
IRAS 00091–0738	3187 + 2306	2005 Jun 30 + 2004 Dec 8	240	240	240	240
IRAS 00456–2904	3187 + 2306	2005 Jul 14 + 2004 Jul 17	240	240	240	240
IRAS 01004–2237	105	2004 Jan 4	120	120	120	120
IRAS 01166–0844	3187 + 2306	2005 Jan 3 + 2005 Jan 15 + 2004 Jul 17	480 ^c	480 ^c	480 ^a	480 ^a
IRAS 01298–0744	105	2005 Jul 14	240	240	120	120
IRAS 01569–2939	2306	2004 Jul 18	240	240	240	240
IRAS 02411+0353	2306	2005 Jan 14	240	240	240	240
IRAS 10190+1322	3187 + 2306	2005 May 22 + 2004 Dec 12	240 + 240 ^d	240 + 240 ^d	480 ^a	480 ^a
IRAS 11387+4116	2306	2005 Jan 11	240	240	240	240
IRAS 11506+1331	3187 + 2306	2005 May 25 + 2005 Jan 3	240	240	240	240
IRAS 13509+0442	2306	2004 Jul 17	240	240	240	240
IRAS 13539+2920	2306	2005 Feb 7	240	240	240	240
IRAS 14060+2919	2306	2004 Jul 16	240	240	240	240
IRAS 15206+3342	105	2004 Jun 24	120	120	180	180
IRAS 15225+2350	2306	2005 Feb 7	240	240	240	240
IRAS 16474+3430	2306	2004 Jul 14	240	240	240	240
IRAS 20414–1651	105	2004 May 14	120	120	120	120
IRAS 21208–0519	3187 + 2306	2004 Nov 13 + 2004 Oct 22	240 ^c	240 ^c	480 ^a	480 ^a
IRAS 22206–2715	3187 + 2306	2004 Nov 15 + 2004 Oct 25	240	240	480 ^a	480 ^a
IRAS 22491–1808	105	2004 Jun 24	120	120	120	120

NOTES.—Col. (1): Object name. Col. (2): PID number: 105 (PI: J. Houck), 2306 (PI: M. Imanishi), and 3187 (PI: S. Veilleux). For sources denoted with “3187 + 2306,” SL and LL spectra were taken by observations with PIDs = 3187 and 2306, respectively. Col. (3): Observing date in UT. Col. (4): Net on-source integration time for SL2 spectroscopy in seconds. Col. (5): Net on-source integration time for SL1 spectroscopy in seconds. Col. (6): Net on-source integration time for LL2 spectroscopy in seconds. Col. (7): Net on-source integration time for LL1 spectroscopy in seconds.

^a Although separate spectroscopy of double nuclei (Kim et al. 2002) was proposed, the same coordinate was observed twice, probably because IRS blue peak up camera (13–18.5 μm) pointed to the same position.

^b We combined SL spectra of double nuclei with $\sim 3''$ separation (Kim et al. 2002), because both nuclei are very faint. The combined SL1 spectrum is smoothly connected to the LL2 spectrum.

^c We combined SL spectra of double nuclei with $\sim 5''$ separation (Kim et al. 2002), because the fainter northern nucleus is very faint. The combined SL1 spectrum is smoothly connected to the LL2 spectrum.

^d SL spectra of the eastern and western nuclei with $\sim 4''$ separation (Kim et al. 2002) are analyzed separately, because both nuclei have similar flux levels. The combined SL1 spectrum is smoothly connected to the LL2 spectrum.

^e Only the brighter northern nucleus (Kim et al. 2002) was analyzed, because (1) the fainter southern nucleus with $\sim 7''$ separation is very faint, (2) it may not be fully covered with the LL spectrum, and (3) the SL1 spectrum of the northern nucleus is smoothly connected to the LL2 spectrum.

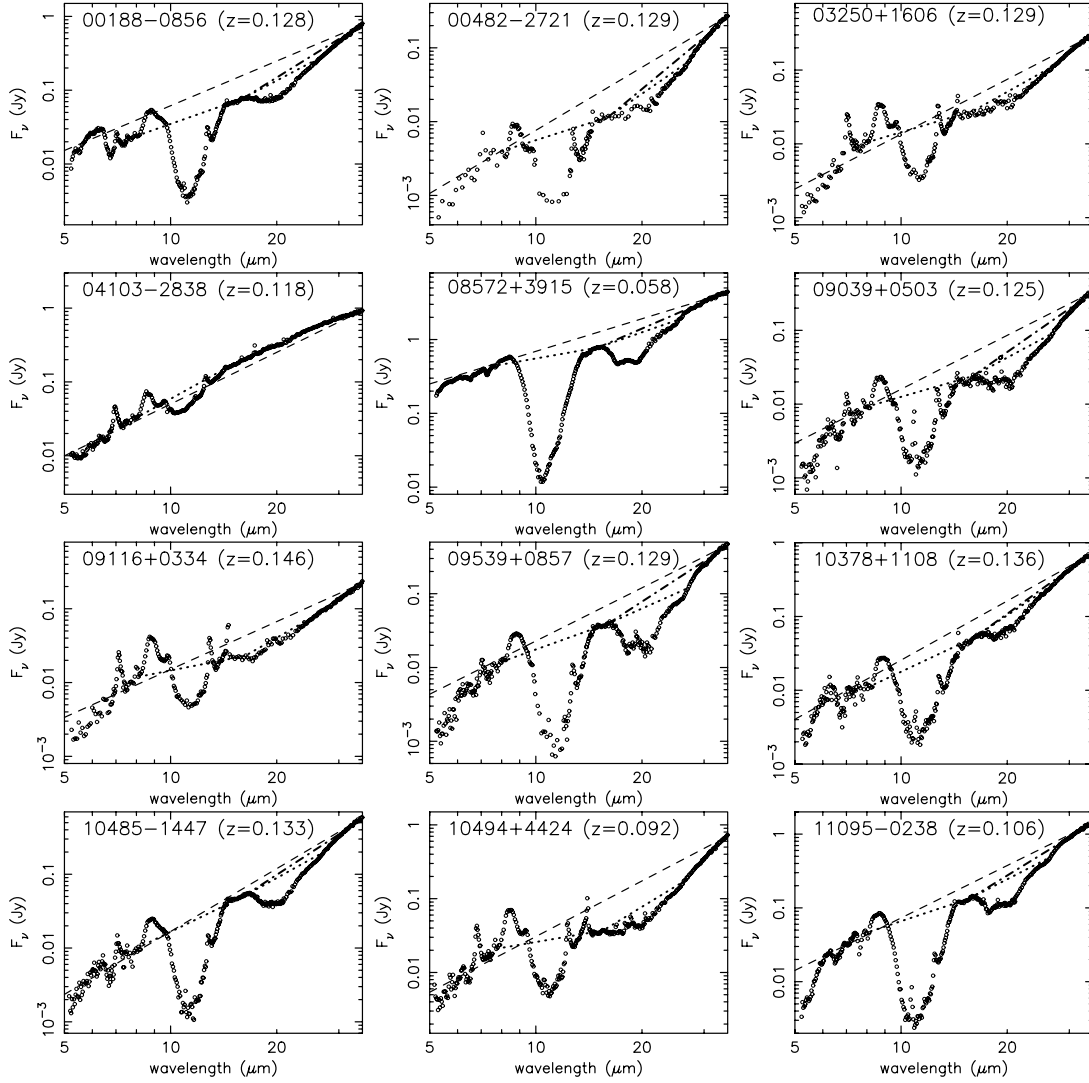


FIG. 3.—Spectra of ULIRGs classified optically as LINERs. The abscissa and ordinate are, respectively, the observed wavelength in μm and flux F_ν in Jy, both shown as decimal logarithmic scale. For all objects, the ratio of the uppermost to lowermost scale in the ordinate is a factor of 1000, to show the variation of the overall spectral energy distribution. *Dashed line*: power-law continuum determined from data points at $\lambda_{\text{rest}} = 5.6$ and $7.1 \mu\text{m}$, and $\lambda_{\text{obs}} = 34\text{--}35 \mu\text{m}$. *Dotted line*: power-law continuum determined from data points at $\lambda_{\text{rest}} = 7.1$ and $14.2 \mu\text{m}$ for $9.7 \mu\text{m}$ silicate absorption, and at $\lambda_{\text{rest}} = 14.2$ and $24 \mu\text{m}$ for $18 \mu\text{m}$ silicate absorption. *Dash-dotted line*: power-law continuum determined from data points at $\lambda_{\text{rest}} = 14.2$ and $29 \mu\text{m}$, used only for ULIRGs showing strong $18 \mu\text{m}$ silicate dust absorption features.

et al. 2002) but are not discussed in detail because no new meaningful information is expected to be obtained.

In summary, the 6.2 and $11.3 \mu\text{m}$ PAH emission features are the main focus of our discussion, and the strength of the $7.7 \mu\text{m}$ PAH feature based on our own definition is shown just for reference. We fit the PAH emission features with Gaussian profiles, which reproduce the observed data reasonably well. The observed rest-frame equivalent widths and luminosities of the 6.2 , 7.7 , and $11.3 \mu\text{m}$ PAH emission features, based on our adopted continuum levels, are summarized in Table 3. The uncertainties coming from the fittings are unlikely to exceed 30%, so once the continuum levels are fixed, the total uncertainties are of this order.

4.2. Silicate Absorption

Based on our adopted continuum levels, the $9.7 \mu\text{m}$ silicate absorption feature is found in almost all ULIRGs, and $18 \mu\text{m}$ silicate absorption is also seen, particularly in ULIRGs with strong $9.7 \mu\text{m}$ absorption (Figs. 3 and 4). Silicate dust grains show two strong absorption peaks centered at $\lambda_{\text{rest}} = 9.7$ and $18.5 \mu\text{m}$. Outside these peaks, a weak absorption tail extends from $\lambda_{\text{rest}} = 8$

to $>30 \mu\text{m}$ (Chiar & Tielens 2006; see also Fig. 6a in this paper). Thus, all wavelengths between $\lambda_{\text{rest}} = 8$ and $30 \mu\text{m}$ show excess absorption, compared to the extrapolation of dust extinction curve from wavelengths outside $\lambda_{\text{rest}} = 8\text{--}30 \mu\text{m}$. Hence, to accurately estimate the optical depths of the $9.7 \mu\text{m}$ ($\tau_{9.7}$) and $18 \mu\text{m}$ (τ_{18}) silicate absorption features, a continuum level must be determined using data at $\lambda_{\text{rest}} < 8$ and $>30 \mu\text{m}$. At $\lambda_{\text{rest}} = 6\text{--}8 \mu\text{m}$, PAH emission can contaminate the continuum level, so that in practice data points at $\lambda_{\text{rest}} \sim 5.6 \mu\text{m}$ are used at the shorter wavelength side. Data points at $\lambda_{\text{rest}} \sim 7.1 \mu\text{m}$ are also less affected by PAH emission and absorption features (Spoon et al. 2006) and so can be used as additional continuum points. Table 4 summarizes the optical depths of the $9.7 \mu\text{m}$ silicate absorption ($\tau_{9.7}$), against power-law continuum levels determined from data points at $\lambda_{\text{rest}} = 5.6$ and $7.1 \mu\text{m}$, and $\lambda_{\text{obs}} = 34\text{--}35 \mu\text{m}$ (the longest part of the IRS spectra) shown as the dashed lines in Figures 3 and 4. A power-law continuum is assumed here, because the intrinsic unabsorbed $5\text{--}35 \mu\text{m}$ spectrum of AGN-heated hot dust emission is roughly approximated by a power-law form (Weedman et al. 2005).

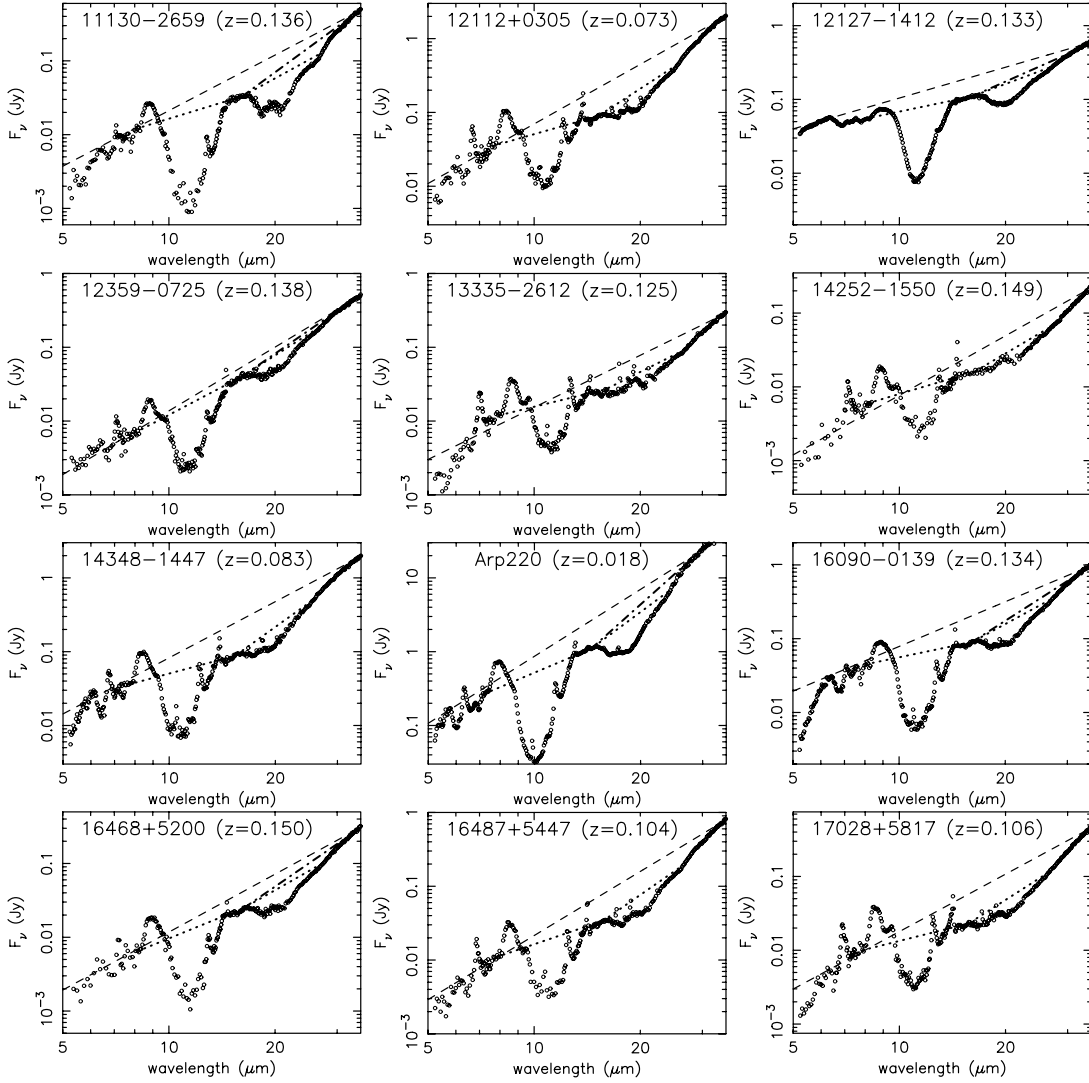


FIG. 3—Continued

The wavelengths of these data points used for the continuum determination are, however, far away from the 9.7 and 18 μm silicate peaks. The optical depth values of the 9.7 and 18 μm silicate absorption features can vary depending on the assumptions of the continuum data points as well as the continuum shape (power-law, linear, spline, or polynomial, etc.). In addition, the power-law continua do not necessarily provide perfect fits to observed data points outside the silicate features. In the spectrum of the weakly obscured ULIRG IRAS 01004–2237 (Fig. 4, *third plot*), data points at $\lambda_{\text{obs}} = 13\text{--}18\ \mu\text{m}$ are clearly above the power-law continuum (*dashed line*), indicating that the power-law continuum only *roughly* represents the overall continuum. The relative ambiguity of the $\tau_{9.7}$ value is small because the 9.7 μm silicate feature is intrinsically strong. However, the optical depth of the much weaker 18 μm silicate absorption feature can be highly uncertain. Consequently, discussions of the presence or absence of strong dust temperature gradients, based on the relative strength of the 9.7 and 18 μm silicate dust absorption features (Fig. 2), could also be highly uncertain, if we use $\tau_{9.7}$ and τ_{18} as defined above.

For this reason, we define $\tau'_{9.7}$ and τ'_{18} . The $\tau'_{9.7}$ value is the optical depth of the 9.7 μm silicate feature against a power-law con-

tinuum determined from data points at $\lambda_{\text{rest}} = 7.1$ and 14.2 μm . The τ'_{18} value is the optical depth of the 18 μm silicate feature against a power-law continuum determined from data points at $\lambda_{\text{rest}} = 14.2$ and 24 μm . These continua are shown as dotted lines in Figures 3 and 4. With these new definitions, the ambiguities are reduced, because the continuum levels are determined using data points just outside the 9.7 and 18 μm features, closer to the absorption peaks. The $\tau'_{9.7}$ values are shown for all ULIRGs in Table 4 (col. [3]). The τ'_{18} values are also shown in Table 4 (col. [4]) for ULIRGs, where the 18 μm feature is clearly seen in absorption.

The τ'_{18} values are appropriate to estimate the strength of the 18 μm silicate feature, if the 18 μm profile is similar to that observed in the ULIRG IRAS 08572+3915 (Spoon et al. 2006). However, the 18 μm silicate profile of the Galactic, highly obscured star GCS3-I extends far beyond $\lambda_{\text{rest}} \sim 24\ \mu\text{m}$ and reaches to $\lambda_{\text{rest}} \sim 29\ \mu\text{m}$ (Chiar & Tielens 2006; see also Fig. 6a in this paper). Thus, we also define τ''_{18} , which is the optical depth of the 18 μm feature against a power-law continuum determined from data points at $\lambda_{\text{rest}} = 14.2$ and 29 μm (shown as dash-dotted lines in Figs. 3 and 4). The τ''_{18} values are also summarized in Table 4 (col. [5]) for ULIRGs with clearly detectable 18 μm silicate absorption.

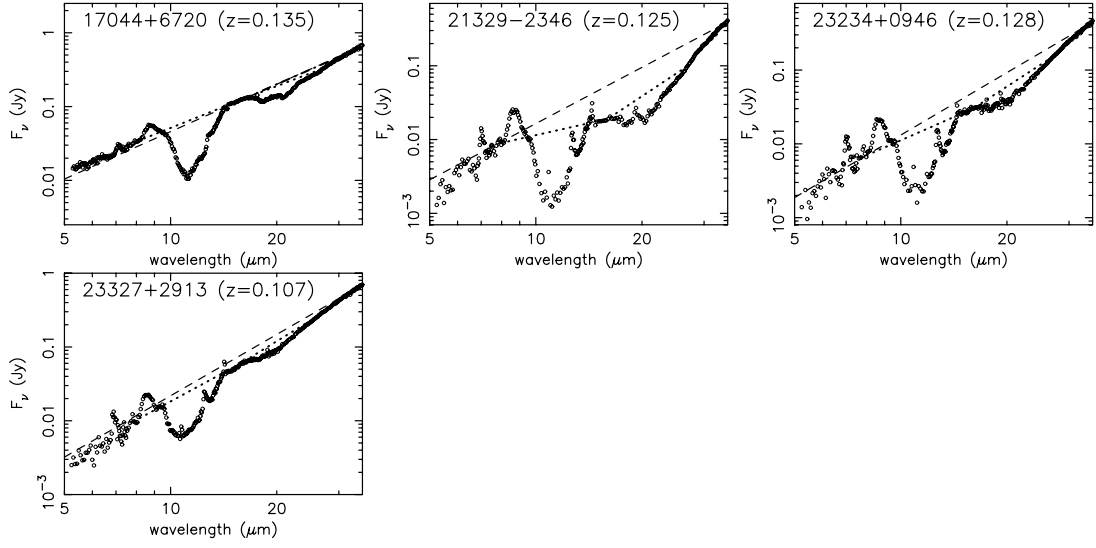


FIG. 3—*Continued*

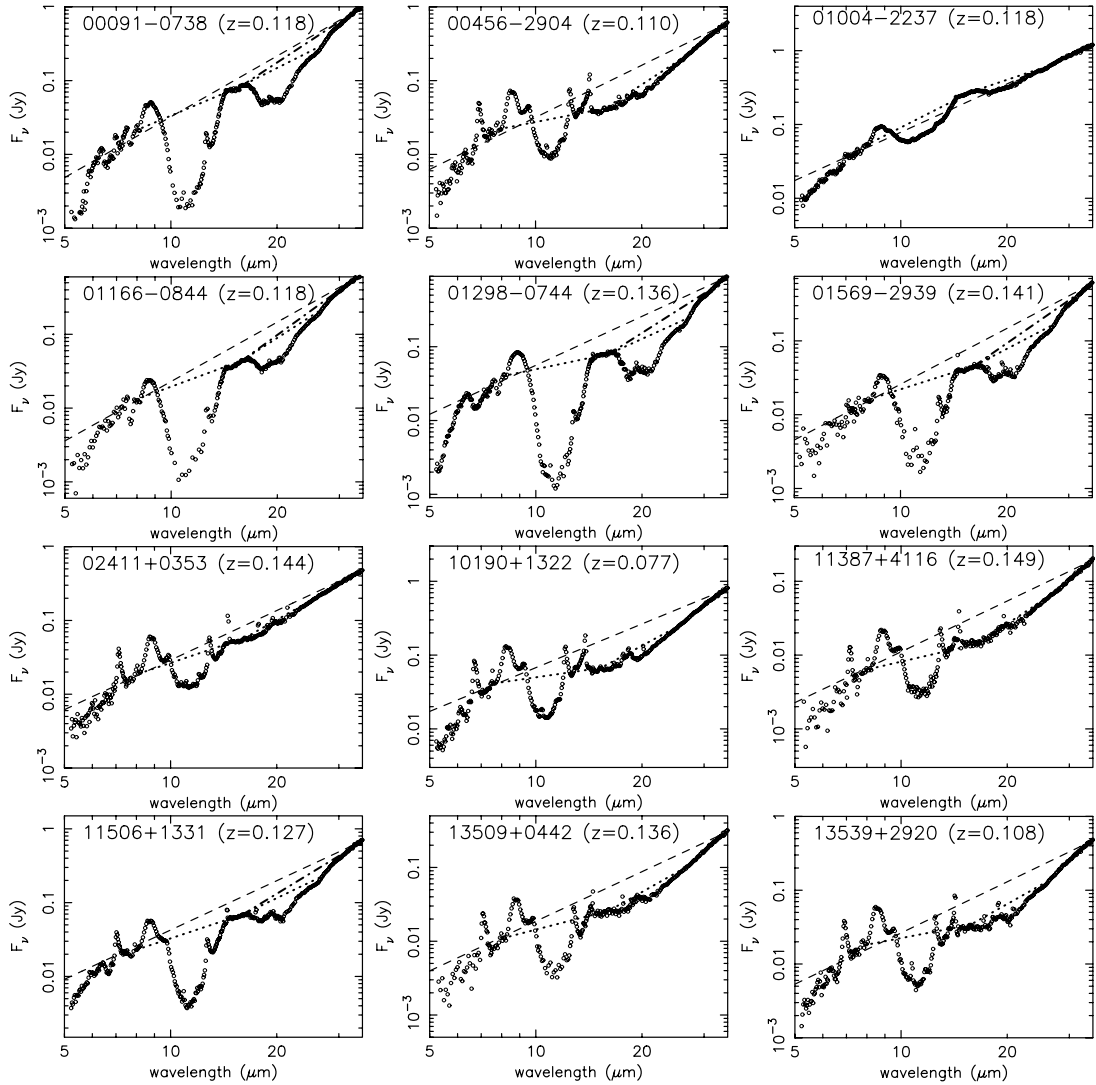
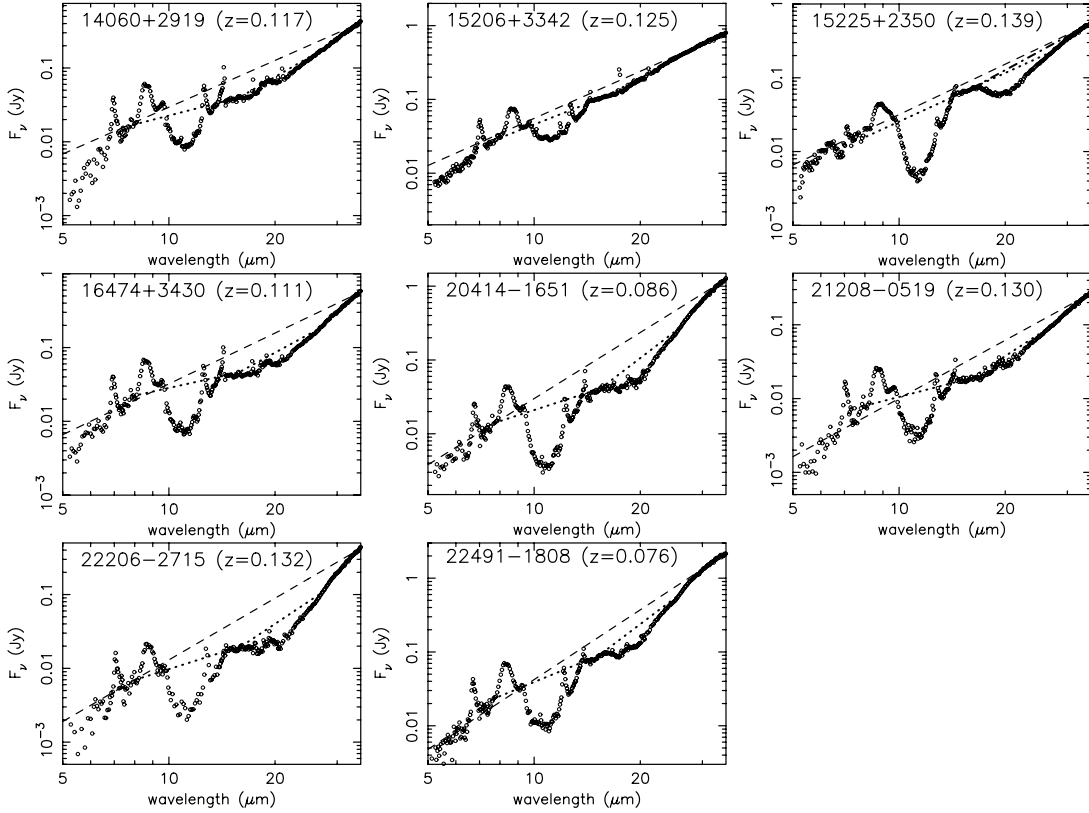


FIG. 4.—Spectra of ULIRGs classified optically as H II regions, shown in the same way as Fig. 3. Symbols are also the same as Fig. 3. For IRAS 10190+1322, SL spectra of the eastern and western nuclei are combined.

FIG. 4—*Continued*

4.3. Ice Absorption

A significant fraction of ULIRGs display dips at the shorter wavelength part of the $6.2 \mu\text{m}$ PAH emission features. To better visualize the dips, Figure 7 presents zoom-in spectra at $\lambda_{\text{obs}} = 5.2\text{--}9 \mu\text{m}$ for ULIRGs showing clear or possible dips. Spectra of two ULIRGs without discernible dips (IRAS 17044+6720 and 00456–2904) are also shown as examples of no clear detection. We ascribe the dips to the $6.0 \mu\text{m}$ H_2O absorption feature (bending mode), as has been found in both Galactic obscured stars (Gibb et al. 2000; Keane et al. 2001) and obscured ULIRGs (Spoon et al. 2002). Since the $6.0 \mu\text{m}$ ice absorption feature is sandwiched by PAH emission features at $\lambda_{\text{rest}} = 5.7$ and $6.2 \mu\text{m}$, it is sometimes difficult to distinguish whether the dip is due to ice absorption or to PAH emission features. For ULIRGs with *clearly* detectable $6.0 \mu\text{m}$ H_2O ice absorption features, their observed optical depths ($\tau_{6.0}$) are summarized in Table 5.

For ULIRGs showing clear ice (and HAC) absorption features, the continuum flux levels just outside of the $6.2 \mu\text{m}$ PAH emission are reduced. For these ULIRGs, approximate $\text{EW}_{6.2\text{PAH}}$ values after these corrections are also shown in Table 3.

5. DISCUSSION

5.1. Magnitudes of Detected Starbursts

PAH emission luminosities can roughly probe the absolute magnitudes of PAH-emitting normal starbursts, if the extinction of the starbursts is say less than about 20–30 mag in A_V , because this amount of A_V produces little flux attenuation (<1 mag) at $\lambda_{\text{rest}} > 5 \mu\text{m}$, aside from the strong $9.7 \mu\text{m}$ silicate absorption feature (Rieke & Lebofsky 1985; Lutz et al. 1996).

Table 3 (cols. [9] and [10]) tabulates the $6.2 \mu\text{m}$ PAH-to-infrared luminosity ratio, $L_{6.2\text{PAH}}/L_{\text{IR}}$, and the $11.3 \mu\text{m}$ PAH-to-infrared luminosity ratio, $L_{11.3\text{PAH}}/L_{\text{IR}}$. The ratios in normal starburst galaxies with

modest dust obscuration ($A_V < 20\text{--}30$ mag) are estimated to be $L_{6.2\text{PAH}}/L_{\text{IR}} \sim 3.4 \times 10^{-3}$ (Peeters et al. 2004) and $L_{11.3\text{PAH}}/L_{\text{IR}} \sim 1.4 \times 10^{-3}$ (Soifer et al. 2002). Figures 8a and 8b display the PAH-to-infrared luminosity ratio (*ordinate*) and PAH equivalent width (*abscissa*) for the 6.2 and $11.3 \mu\text{m}$ PAH features, respectively. The observed $L_{6.2\text{PAH}}/L_{\text{IR}}$ and $L_{11.3\text{PAH}}/L_{\text{IR}}$ ratios in the majority of the observed ULIRGs are systematically lower than those seen in modestly obscured normal starburst galaxies, suggesting that the detected starburst activity as measured by the PAH emission in ULIRGs is energetically insignificant. The dominant energy sources of ULIRGs must therefore be either (1) highly obscured ($A_V \gg 30$ mag) starbursts, where the emitted PAH luminosities are substantially attenuated by dust extinction, or (2) buried AGNs, which produce intrinsically weak PAH emission.

5.2. Candidates of ULIRGs Possessing Powerful Buried AGNs

5.2.1. Low Equivalent Widths of PAH Emission Features

Whether the dominant energy sources of ULIRGs' cores are highly obscured normal starbursts with mixed dust/source geometry (Fig. 1a) or buried AGNs with centrally concentrated energy source geometry (Figs. 1b and 1c) can be distinguished using the equivalent width of the PAH emission. Since the PAH equivalent width (EW_{PAH}) must always be large in a normal starburst regardless of the amount of dust extinction, a small EW_{PAH} value requires a contribution from a PAH-free continuum-emitting energy source (§ 1).

Figures 9a, 9b, and 9c compare the equivalent widths of the 6.2 , 7.7 , and $11.3 \mu\text{m}$ PAH emission features. Positive correlations are found in all plots, suggesting that (1) variation of the relative PAH emission strength is modest, with some level of scatter and (2) PAH-emitting starbursts usually show strong PAH features in all the emission modes at 6.2 , 7.7 , and $11.3 \mu\text{m}$.

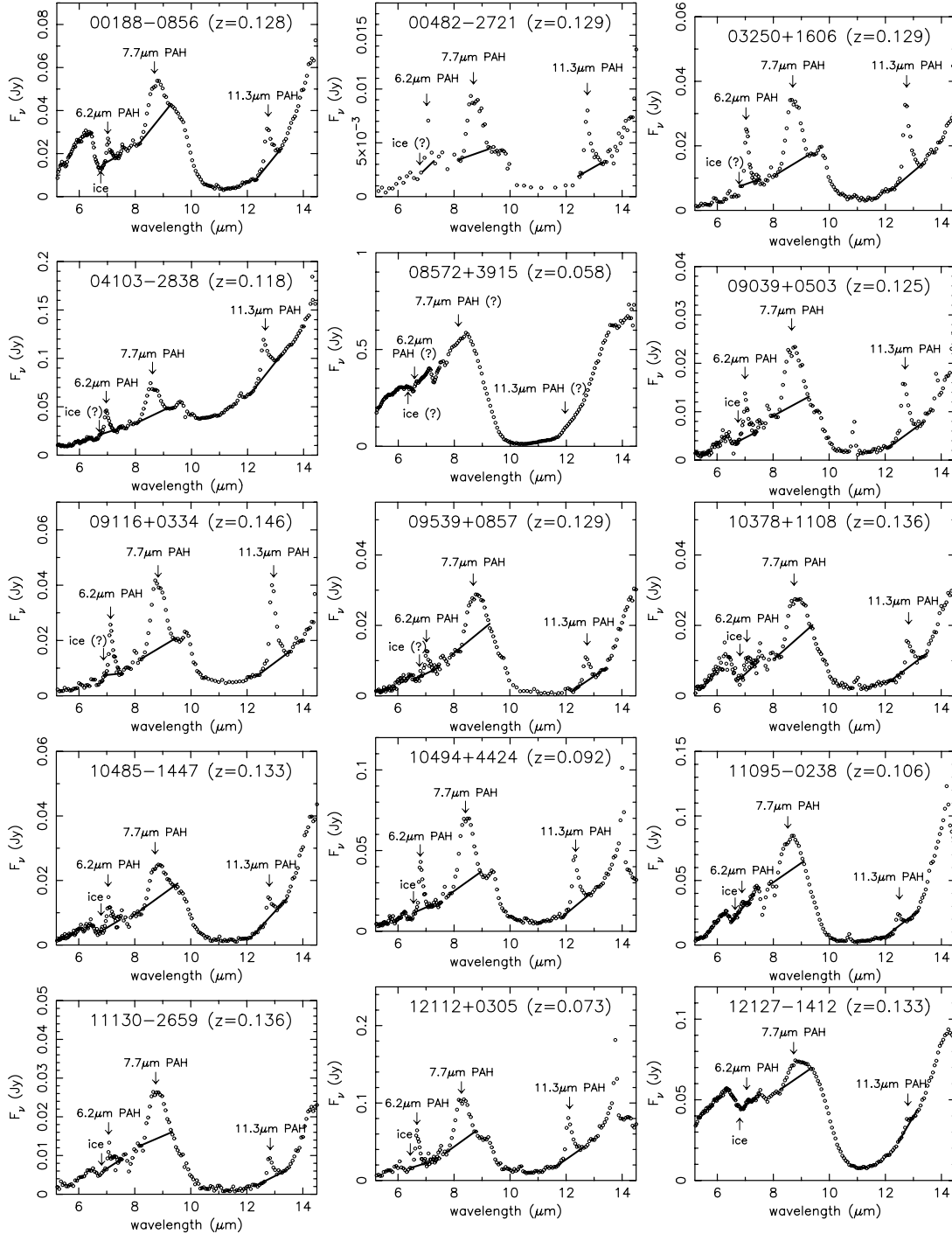


FIG. 5.—Spectra of all ULIRGs at $\lambda_{\text{obs}} = 5.2\text{--}14.5\ \mu\text{m}$, to zoom in on the PAH emission features and display them in detail. The abscissa and ordinate are, respectively, observed wavelength in μm and flux in Jy, both shown at linear scale. For IRAS 10190+1322, the spectra of the eastern and western nuclei are shown separately. The expected wavelengths of the 6.2, 7.7, and 11.3 μm PAH emission features, as well as 6.0 μm H₂O ice absorption feature, are indicated as down arrows. A question mark is added when detection is unclear. The solid lines are adopted continuum levels to estimate the strengths of the detected PAH emission features.

Brandl et al. (2006) have investigated the 6.2 and 11.3 μm PAH emission equivalent widths, based on similar definitions of PAH strength to ours, in nearby, normal starburst galaxies without obvious Seyfert signatures in the optical (18 sources), and found that $\text{EW}_{6.2\text{PAH}} = 561 \pm 73\ \text{nm}$ (median 553 nm) and $\text{EW}_{11.3\text{PAH}} = 617 \pm 187\ \text{nm}$ (median 597 nm). The scatter is larger for the 11.3 μm PAH emission feature than the 6.2 μm PAH feature. No starburst galaxies show values smaller than $\text{EW}_{6.2\text{PAH}} = 459\ \text{nm}$ and $\text{EW}_{11.3\text{PAH}} = 316\ \text{nm}$. Since the PAH equivalent

widths are independent of the infrared luminosities of the starburst galaxies in the range $L_{\text{IR}} = 10^{9.7}\text{--}10^{11.6}\ L_{\odot}$ (Brandl et al. 2006), we apply these values to our ULIRG sample ($L_{\text{IR}} > 10^{12}\ L_{\odot}$).

The median values for the 6.2 μm PAH emission equivalent widths (rest-frame) are $\text{EW}_{6.2\text{PAH}} = 230\ \text{nm}$ for the whole of our ULIRG sample (LINER + H II region), 225 nm for LINER ULIRGs, and 300 nm for H II region ULIRGs. The median 11.3 μm PAH equivalent widths are $\text{EW}_{11.3\text{PAH}} = 390\ \text{nm}$ (LINER + H II region), 350 nm (LINER), and 460 nm (H II region). The median

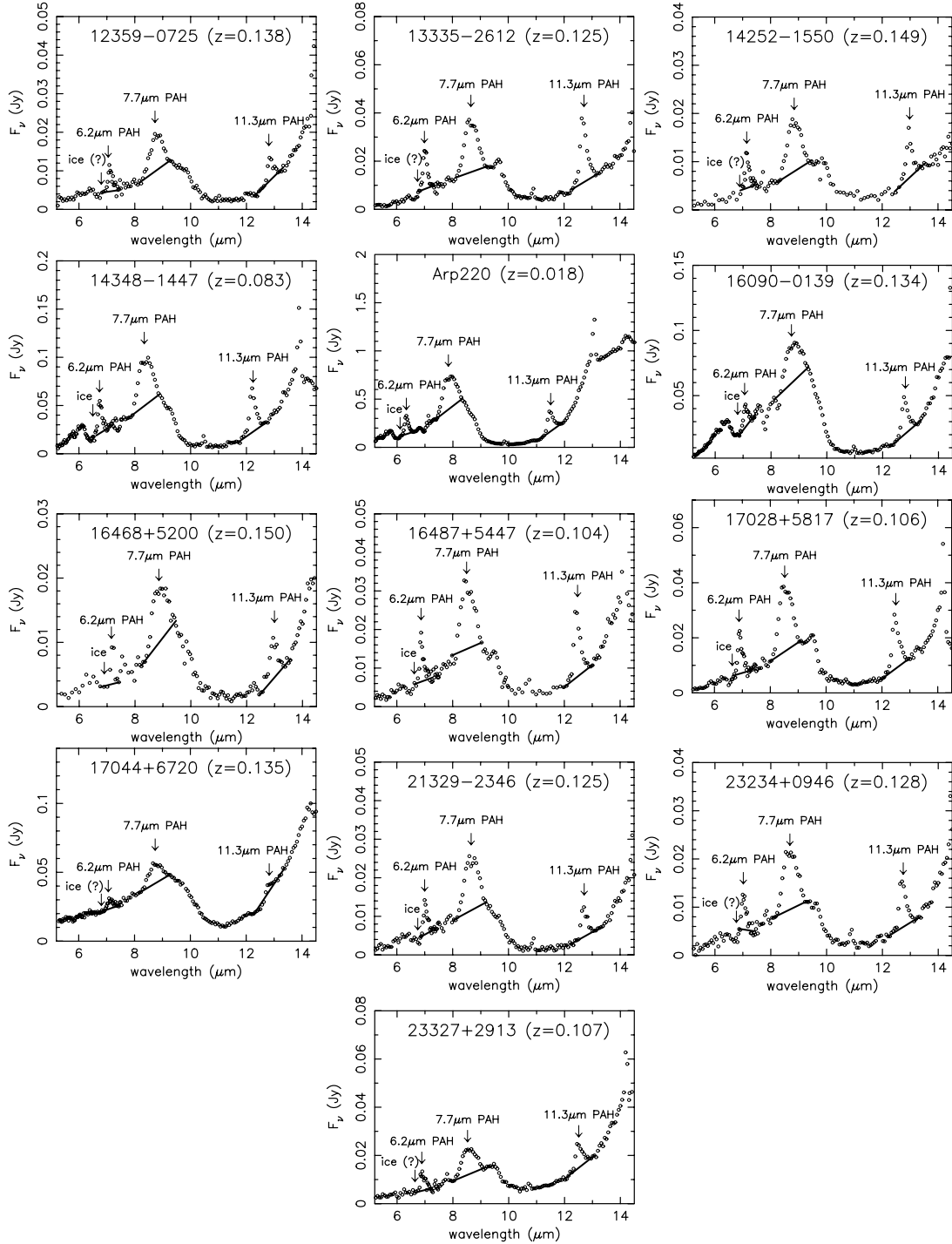


FIG. 5—Continued

equivalent widths of both PAH features are smaller in ULIRGs than in starburst galaxies (Brandl et al. 2006), suggesting that PAH-free continuum emission, originating in AGN-heated hot dust, contributes more strongly to the observed fluxes in ULIRGs. We assume that if a PAH equivalent width is more than a factor of ~ 3 smaller than the median value in normal starburst galaxies, then an AGN-originating PAH-free continuum contributes importantly to the observed 5–12 μm flux and dilutes the PAH emission. With this criterion, ULIRGs with $\text{EW}_{6.2\text{PAH}} < 180 \text{ nm}$ and $\text{EW}_{11.3\text{PAH}} < 200 \text{ nm}$ are classified as sources displaying clear signatures of luminous AGNs. Since these adopted values are clearly below the minimum values observed in normal starburst

galaxies, we believe the criterion is reasonable and conservative to pick up luminous buried AGN candidates.

For the 7.7 μm PAH equivalent widths, based on our definition of the PAH strength (§ 4.1), the median 7.7 μm PAH equivalent widths in our ULIRG sample are $\text{EW}_{7.7\text{PAH}} = 490 \text{ nm}$ (LINER) and 695 nm (H II region). In our ULIRG sample, the median PAH equivalent widths are smaller in LINER ULIRGs than in H II region ULIRGs for all of the 6.2, 7.7, and 11.3 μm features. We classify ULIRGs with $\text{EW}_{7.7\text{PAH}} < 230 \text{ nm}$, one-third of the median value for H II region ULIRGs, as those showing luminous buried AGN signatures. Since some fraction of H II region ULIRGs may contain luminous buried AGNs, this

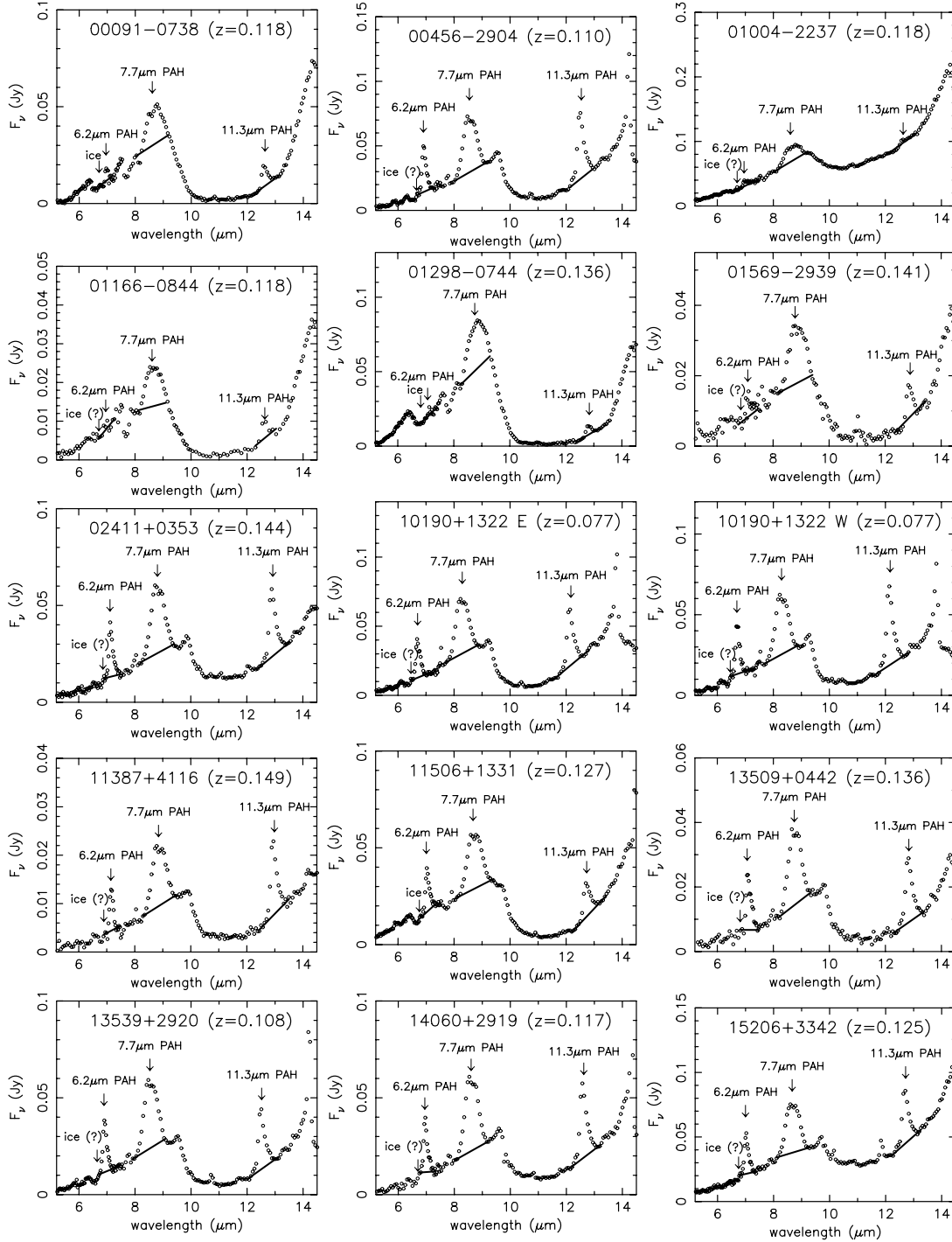


FIG. 5—Continued

$EW_{7.7PAH}$ threshold is even more conservative than those for the 6.2 and 11.3 μm PAH emission features, as determined above.

Table 6 (cols. [2]–[4]) presents detection or nondetection of buried AGN signatures based on the PAH equivalent width threshold. The fraction of ULIRGs with clearly detectable buried AGN signatures is 16/48 (33%; LINER + H II region), 10/28 (36%; LINER) and 6/20 (30%; H II region) based on observed $EW_{6.2PAH}$, or 20/48 (42%; LINER + H II region), 13/28 (46%; LINER), and 7/20 (35%; H II region) based on $EW_{6.2PAH}$ after correction for the ice and HAC absorption feature. The fraction is 4/48 (8%; LINER + H II region), 3/28 (11%; LINER), and 1/20 (5%; H II region) based on observed $EW_{7.7PAH}$, 12/48 (25%;

LINER + H II region), 8/28 (29%; LINER), and 4/20 (20%; H II region) based on observed $EW_{11.3PAH}$. The fraction of ULIRGs where very low PAH equivalent widths are found in any of the features is 19/48 (40%; LINER + H II region), 13/28 (46%; LINER), and 6/20 (30%; H II region). If the corrected $EW_{6.2PAH}$ values are adopted, it is 22/48 (46%; LINER + H II region), 15/28 (54%; LINER), and 7/20 (35%; H II region).

Since we adopted a conservative threshold, this method provides only candidate ULIRGs where PAH-free continuum emission from buried AGNs contribute very importantly to the observed IRS spectra. In starburst/buried-AGN composite ULIRGs (Fig. 1c), PAH emission from starbursts surrounding the central buried

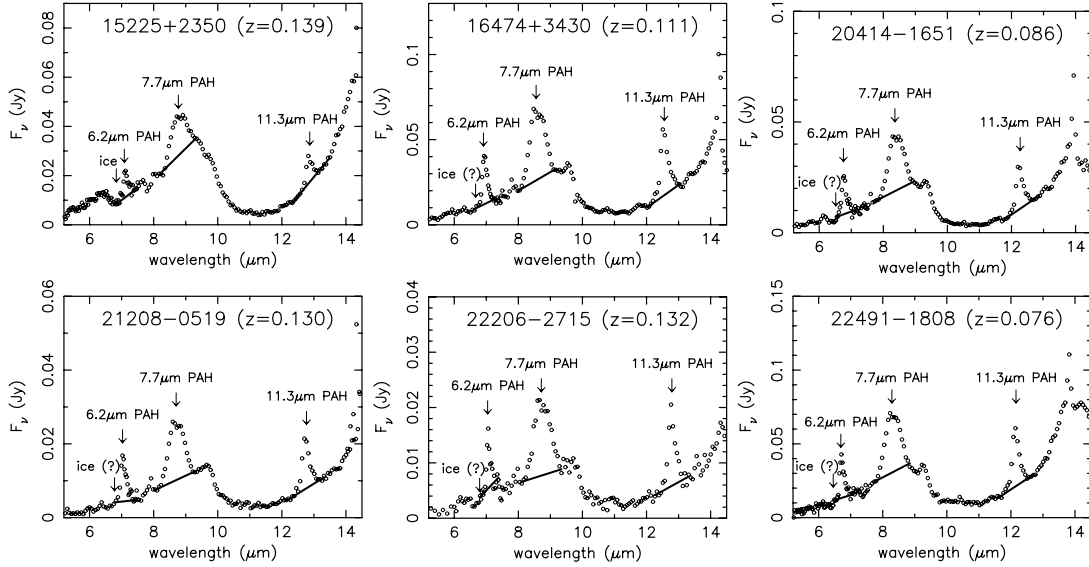


FIG. 5—Continued

AGNs is less obscured and less flux-attenuated, making the observed PAH equivalent widths relatively large, particularly when AGNs are deeply buried. Such buried AGNs should be searched for using other methods.

5.2.2. Absolute Optical Depths of Dust Absorption Features

In a normal starburst with mixed dust/source geometry (Fig. 1a), the foreground, less-obscured, less-attenuated emission (which shows only weak dust absorption features) dominates the observed flux, and therefore the observed optical depths of dust absorption features cannot exceed a certain threshold, unless very unusual dust composition patterns are assumed (Imanishi & Maloney 2003; Imanishi et al. 2006a). In a buried AGN, with centrally concentrated source geometry (Fig. 1b), the so-called foreground screen dust model is applicable, and the observed optical depths can be arbitrarily large. Hence, detection of strong dust absorption features, with optical depths substantially larger than the upper limit set by the mixed dust/source geometry, argues for a foreground screen dust geometry, as expected from a buried AGN (Imanishi & Maloney 2003; Imanishi et al. 2006a).

To differentiate the geometry using the optical depths of dust absorption features, the $9.7 \mu\text{m}$ silicate dust absorption is particularly useful, because (1) it is much stronger than the $18 \mu\text{m}$ silicate feature (Fig. 6a) and (2) its intrinsic strength has been widely investigated in Galactic obscured objects (Roche & Aitken 1984, 1985). Assuming the Galactic dust extinction curve, the upper limit for the optical depth of the $9.7 \mu\text{m}$ silicate absorption feature in the mixed dust/source geometry can be derived. The Appendix presents detailed calculations and we obtain a maximum value of $\tau'_{9.7} < 1.7$ for mixed dust/source geometry. In fact, none of the starburst galaxies ($L_{\text{IR}} < 10^{11.6} L_{\odot}$) lacking optical Seyfert signatures so far observed with *Spitzer* IRS (Brandt et al. 2006) show values that exceed the threshold ($\tau'_{9.7} = 1.7$). Thus, it is reasonable to assume that ULIRGs with $\tau'_{9.7}$ substantially larger than 1.7 possess buried AGNs with centrally concentrated energy source geometry.

In ULIRGs, it is possible that the dust extinction curve is very different from the Galactic one, because dust coagulation in the high-density environment in ULIRGs' cores could increase typical dust grain size up to a few microns (Laor & Draine 1993; Maiolino et al. 2001a, 2001b; Imanishi 2001). However, we be-

lieve this possibility has minimal effect on our conclusions about the geometry, as long as we adopt $\tau'_{9.7}$ (peak optical depth of the $9.7 \mu\text{m}$ feature, relative to the nearby continuum just outside the feature; see § 4.2), rather than $\tau_{9.7}$ (peak optical depth of the $9.7 \mu\text{m}$ feature and continuum, relative to continuum outside $8-30 \mu\text{m}$; see § 4.2). Figure 6b shows the absorption optical depth of astronomical silicate dust grains of three different sizes. The absolute absorption optical depth at the peak of the $9.7 \mu\text{m}$ feature becomes stronger with increasing dust grain size. However, both the absorption optical depth inside the $9.7 \mu\text{m}$ feature and the nearby continuum just outside the feature increase in a very similar way with increasing dust size, so that the strength of $9.7 \mu\text{m}$ feature relative to the nearby continuum ($\tau'_{9.7}$), is virtually unchanged with varying dust grain size. The argument of the geometry based on $\tau'_{9.7}$ values are therefore robust to possible increase of dust grain size in ULIRGs' cores. We conservatively take ULIRGs with $\tau'_{9.7} > 2$ as candidates for harboring luminous centrally concentrated buried AGNs.

Individual ULIRGs with clearly detectable buried AGN signatures based on large $\tau'_{9.7}$ values (> 2 ; Table 4) are marked with open circles in Table 6 (col. [5]). The median $\tau'_{9.7}$ values are 1.8 (LINER + H II region), 2.1 (LINER), and 1.6 (H II region). The fraction of ULIRGs with $\tau'_{9.7} > 2$, indicative of buried AGNs with centrally concentrated energy source geometry, is 20/48 (42%; LINER + H II region), 14/28 (50%; LINER), and 6/20 (30%; H II region). The fraction of ULIRGs showing buried AGN signatures is again higher in LINER ULIRGs than H II region ULIRGs, as was previously found based on the PAH-equivalent-width method (§ 5.2.1).

We must raise three cautions regarding this method. First, a $\tau'_{9.7}$ value smaller than 1.7 does not necessarily preclude the presence of a luminous AGN, because a weakly obscured AGN, of course, shows a small $\tau'_{9.7}$ value. The best example is IRAS 01004-2237 ($\tau_{9.7} \sim 0.4$), which shows very low PAH emission equivalent widths, as expected from an AGN-dominated source.

Next, we used the threshold of $\tau'_{9.7} > 2$, based on the *observed* values. For ULIRGs with detectable PAH emission from modestly obscured starbursts, the observed fluxes are a superposition of the emission from the highly obscured cores and the starburst emission (Fig. 1c), and so the *observed* $\tau'_{9.7}$ values are reduced

TABLE 3
PAH EMISSION

Object (1)	EW _{6.2PAH} (nm) (2)	EW _{7.7PAH} ^a (nm) (3)	EW _{11.3PAH} (nm) (4)	EW _{3.3PAH} (nm) (5)	$L_{6.2PAH}$ (10^{42} ergs s ⁻¹) (6)	$L_{7.7PAH}$ ^a (10^{42} ergs s ⁻¹) (7)	$L_{11.3PAH}$ (10^{42} ergs s ⁻¹) (8)	$L_{6.2PAH}/L_{IR}$ (10^{-3}) (9)	$L_{11.3PAH}/L_{IR}$ (10^{-3}) (10)
IRAS 00188–0856	85 (~40)	305	290	50	3.5	17	3.4	0.4	0.4
IRAS 00482–2721	260	710	465	...	1.8	4.8	0.9	0.5	0.2
IRAS 03250+1606	325	740	570	80	7.0	18	4.6	1.6	1.1
IRAS 04103–2838	165	365	120	...	8.4	21	6.6	1.4	1.1
IRAS 08572+3915	<15	<100	<30	<5	<1.9	<18	<0.5	<0.4	<0.1
IRAS 09039+0503	325 (~200)	625	580	95	3.7	11	2.2	0.8	0.5
IRAS 09116+0334	360	730	630	75	9.4	27	7.4	1.9	1.5
IRAS 09539+0857	145	405	390	65	2.5	12	1.3	0.6	0.3
IRAS 10378+1108	110 (~50)	450	210	40	2.2	14	1.5	0.3	0.2
IRAS 10485–1447	225 (~150)	400	150	80	3.7	11	1.1	0.6	0.2
IRAS 10494+4424	345 (~250)	715	495	110	6.1	19	3.1	1.1	0.6
IRAS 11095–0238	20 (~10)	250	110	150	1.0	17	0.9	0.2	0.1
IRAS 11130–2659	95 (~60)	490	280	...	1.9	13	1.0	0.4	0.2
IRAS 12112+0305	390	565	425	100+160	6.1	15	3.2	0.8	0.4
IRAS 12127–1412	5 (~3)	95	20	0	0.7	11	0.6	0.1	0.1
IRAS 12359–0725	220	445	165	75	3.1	8.8	1.1	0.6	0.2
IRAS 13335–2612	310	700	610	...	6.6	17	5.1	1.5	1.2
IRAS 14252–1550	225	685	320	60+60	3.6	12	2.3	0.6	0.4
IRAS 14348–1447	260 (~110)	540	500	70+110	5.9	19	3.6	0.8	0.5
Arp 220	215 (~120)	410	315	80	1.5	5.1	0.8	0.3	0.2
IRAS 16090–0139	110 (~70)	265	325	75	8.1	30	5.6	0.7	0.5
IRAS 16468+5200	230	470	435	120	2.9	11	2.0	0.7	0.5
IRAS 16487+5447	250	580	570	70	2.8	9.6	2.2	0.5	0.4
IRAS 17028+5817	375	800	595	120	4.6	14	2.8	0.9	0.6
IRAS 17044+6720	55	165	30	10	3.9	13	1.0	0.8	0.2
IRAS 21329–2346	230	635	370	50	3.0	12	1.3	0.6	0.3
IRAS 23234+0946	220	695	445	75	3.1	11	2.0	0.7	0.4
IRAS 23327+2913	255	485	175	45	2.5	6.9	1.4	0.6	0.3
IRAS 00091–0738	50 (~30)	350	255	...	1.3	15	1.5	0.2	0.3
IRAS 00456–2904	385	750	545	...	11	27	7.6	1.9	1.4
IRAS 01004–2237	45	190	8	...	3.1	19	0.5	0.4	0.1
IRAS 01166–0844	30	430	190	...	0.5	8.5	0.7	0.1	0.2
IRAS 01298–0744	10 (~6)	380	155	...	0.7	37	1.1	0.1	0.1
IRAS 01569–2939	115	555	255	...	3.1	20	2.0	0.5	0.3
IRAS 02411+0353	300	695	375	...	13	38	7.9	2.1	1.3
IRAS 10190+1322E	375	680	510	95	4.4	12	3.0	1.2	0.8
IRAS 10190+1322W	375	745	685	50	4.6	11	3.6	1.2	0.9
IRAS 11387+4116	265	680	475	75	4.2	15	3.9	0.7	0.7
IRAS 11506+1331	195 (~120)	540	305	95	8.5	26	3.4	1.1	0.4
IRAS 13509+0442	490	795	590	135	9.7	22	4.7	1.5	0.7
IRAS 13539+2920	355	755	600	85	7.6	22	4.4	1.9	1.1
IRAS 14060+2919	340	860	530	150	8.6	27	6.5	2.0	1.5
IRAS 15206+3342	225	525	230	55	12	32	7.7	2.0	1.3
IRAS 15225+2350	145 (~100)	290	130	40	5.0	16	2.0	1.0	0.4
IRAS 16474+3430	390	820	575	105	9.2	27	5.8	1.7	1.1
IRAS 20414–1651	330	705	465	75	3.2	10	1.7	0.5	0.3
IRAS 21208–0519	505	810	420	100	6.0	14	2.7	1.5	0.7
IRAS 22206–2715	210	975	565	...	3.1	13	2.7	0.6	0.5
IRAS 22491–1808	325	695	460	...	3.9	12	2.6	0.8	0.5

NOTES.—Col. (1): Object name. Col. (2): Rest-frame equivalent width of the 6.2 μ m PAH emission. The values in parentheses are approximate ones after correction for the ice and HAC absorption features. Col. (3): Rest-frame equivalent width of the 7.7 μ m PAH emission. Col. (4): Rest-frame equivalent width of the 11.3 μ m PAH emission. Col. (5): Rest-frame equivalent width of the 3.3 μ m PAH emission, taken from Imanishi et al. (2006a). Ellipses signify no data. Col. (6): Luminosity of the 6.2 μ m PAH emission in units of 10^{42} ergs s⁻¹. Col. (7): Luminosity of the 7.7 μ m PAH emission in units of 10^{42} ergs s⁻¹. Col. (8): Luminosity of the 11.3 μ m PAH emission in units of 10^{42} ergs s⁻¹. Col. (9): The 6.2 μ m PAH-to-infrared luminosity ratio in units of 10^{-3} . The ratio for normal starbursts with modest dust obscuration ($A_V < 20$ –30 mag) is $\sim 3.4 \times 10^{-3}$ (Peeters et al. 2004). Col. (10): The 11.3 μ m PAH-to-infrared luminosity ratio in units of 10^{-3} . The ratio for normal starbursts with modest dust obscuration ($A_V < 20$ –30 mag) is $\sim 1.4 \times 10^{-3}$ (Soifer et al. 2002).

^a We regard flux excess at $\lambda_{\text{rest}} = 7.3$ –8.1 μ m above an adopted continuum level as 7.7 μ m PAH emission, to reduce the effects of the strong 9.7 μ m silicate dust absorption feature. Our definition is different from those in many previous literatures.

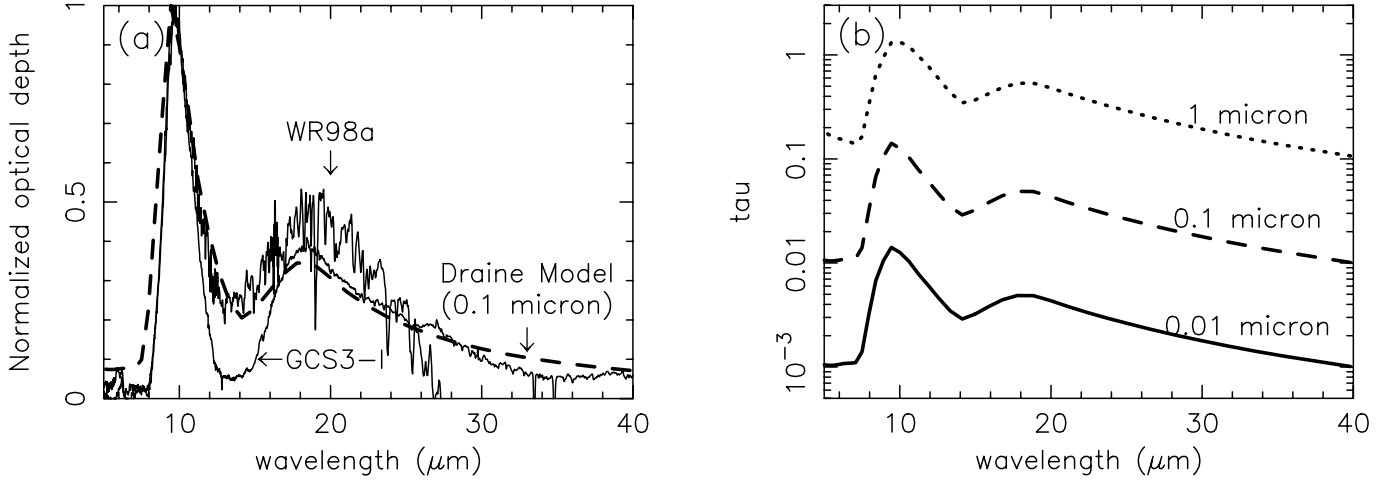


FIG. 6.—Silicate dust absorption profile. (a) Solid lines are the observed silicate absorption profiles of the Galactic stars GCS 3-I and WR 98a (Chiar & Tielens 2006). The dashed line is the theoretical absorption profile of “astronomical silicate” grains with 0.1 μm in size (Draine & Lee 1984). All profiles are normalized at the 9.7 μm absorption peak. The silicate absorption profile extends from 8 to >30 μm . (b) Theoretical absorption profile of “astronomical silicate” grains (Draine & Lee 1984) of different sizes, 1 μm (dotted line), 0.1 μm (dashed line), and 0.01 μm (solid line). Adopted from <http://www.astro.princeton.edu/~draine/dust/dust.diel.html>. The ordinate is in arbitrary units but is proportional to the actual optical depth. In this figure, only absorption is taken into account, and scattering is neglected, because scattering is negligible in this wavelength range (see the above Web page).

compared to the $\tau'_{9.7}$ values *only toward the highly obscured energy sources at ULIRGs' cores*. Since subtraction of the modestly obscured starburst component, using a starburst template spectrum, introduces an additional uncertainty, we do not attempt it. Thus, we may miss ULIRGs that show large $\tau'_{9.7}$ values (>2) *toward the cores* but small *observed* $\tau'_{9.7}$ values (<2) due to the starburst dilution.

Finally, a starburst nucleus obscured by foreground absorbing dust (Fig. 1d) can also show a $\tau'_{9.7}$ value larger than the threshold determined by the mixed dust/source geometry, because in this case, a foreground screen dust model is applicable. Assuming Galactic dust properties, $\tau'_{9.7} > 2$ corresponds to $A_V > 20$ –40 mag (Roche & Aitken 1984, 1985). If ULIRGs are the mergers of gas-rich spiral galaxies (Sanders et al. 1988a), gas and dust in the original galaxies can be quickly transferred to the nuclear regions (Barnes & Hernquist 1996; Mihos & Hernquist 1996). The bulk of possibly newborn stars and resulting dust are also expected to distribute in the nuclear regions. In fact, most of high-density gas, dust, and energy sources (starbursts and/or AGNs) in ULIRGs are observed to be concentrated to nuclear cores with $< \text{kpc}$ scale (Downes & Solomon 1998; Soifer et al. 2000; Imanishi et al. 2006b). The amount of gas/dust in ULIRGs' hosts at greater than kpc scale (which are morphologically disturbed) is unlikely to substantially exceed that of the original merging gas/dust-rich galaxies. In our Galaxy, a gas-rich spiral, the A_V value is estimated to be ≥ 30 mag in an equatorial direction (Rieke et al. 1989) and should be much smaller in other directions. It is thus not easy to produce such a large A_V value (>20 –40 mag) with dust in a host galaxy, unless the host is viewed from a direction with particularly large dust column density. Although a small fraction of ULIRGs with $\tau'_{9.7} > 2$ may have this geometry (Fig. 1d), it is very unlikely that the majority do.

5.2.3. Strong Dust Temperature Gradients

A buried AGN with centrally concentrated energy source geometry (Fig. 1b) should show a strong dust temperature gradient (Fig. 2), whereas a normal starburst nucleus with mixed dust/source geometry does not, whether it is obscured by foreground dust (Fig. 1d) or not (Fig. 1a). The presence of a strong dust temperature gradient thus supports the buried AGN scenario,

and it can be detected by comparing the optical depths of dust absorption features at different infrared wavelengths (Dudley & Wynn-Williams 1997; Imanishi 2000, 2006a; Imanishi & Maloney 2003). As Figure 2 illustrates, in a buried AGN, the dust column density estimated from observations at $\lambda_{\text{rest}} > 3$ μm probes deeper inside of the obscuring material as the wavelength is decreased [i.e., $A_V(3 \mu\text{m}) > A_V(10 \mu\text{m}) > A_V(20 \mu\text{m})$]. In this geometry, dust with a fixed temperature range from a fixed layer contributes strongly to the observed continuum emission at each wavelength, with a small contribution from other components (Fig. 2). Thus, in a starburst/buried AGN composite (Fig. 1c), PAH-free continuum emission from the AGN-heated dust should be detectable even if its optical depth is larger than unity at a particular observed wavelength, provided the buried AGN is intrinsically sufficiently luminous that the flux-attenuated ($e^{-\tau}$; $\tau > 1$) AGN-heated dust emission contributes significantly to the observed flux. This has been demonstrated in the 3–4 μm spectra of the ULIRGs IRAS 08572+3915 and UGC 5101 ($\tau_{3-4 \mu\text{m}} > 5$; Imanishi 2000, 2001; Imanishi & Maloney 2003).

On the other hand, in a normal starburst (Fig. 1a), no significant dust temperature gradient is expected, and so the estimated dust extinction should be similar at different wavelengths [$A_V(3 \mu\text{m}) \sim A_V(10 \mu\text{m}) \sim A_V(20 \mu\text{m})$]. Since in a normal starburst, emission arises uniformly within an emitting volume, we can typically probe emission with optical depth less than about unity. If the extinction is large enough that the short wavelengths do not probe the whole emitting region, the estimated dust extinction would decrease with decreasing wavelength [$A_V(3 \mu\text{m}) < A_V(10 \mu\text{m}) < A_V(20 \mu\text{m})$]. Hence, the relations among $A_V(3 \mu\text{m})$, $A_V(10 \mu\text{m})$, and $A_V(20 \mu\text{m})$ are totally different between a buried AGN (Fig. 1b) and a normal starburst (Fig. 1a).

For a buried AGN, the $\tau'_{18}/\tau'_{9.7}$ or $\tau''_{18}/\tau'_{9.7}$ ratio in the *Spitzer* IRS spectrum becomes small compared to the value in the absence of a strong dust temperature gradient. The $\tau'_{18}/\tau'_{9.7}$ ($\tau''_{18}/\tau'_{9.7}$) ratio is also insensitive to possible changes in dust grain size in ULIRGs, because the relative strength of τ'_{18} (τ''_{18}) and $\tau'_{9.7}$ varies little with varying dust grain size (Fig. 6b). Furthermore, since both absorption features are due to silicate dust grains, the ratio is unaffected by possibly different dust composition patterns in individual galaxies.

TABLE 4
OPTICAL DEPTH OF THE 9.7 μm AND 18 μm SILICATE DUST ABSORPTION FEATURES

Object (1)	$\tau_{9.7}$ (2)	$\tau'_{9.7}$ (3)	τ'_{18} (4)	τ''_{18} (5)	$\tau'_{18}/\tau'_{9.7}$ (6)	$\tau''_{18}/\tau'_{9.7}$ (7)
IRAS 00188–0856	3.1	2.5	0.55	0.65	0.22	0.26
IRAS 00482–2721	2.6	2.1	0.6	0.75	0.29	0.36
IRAS 03250+1606	1.8	1.6	...	...	...	...
IRAS 04103–2838	0.5	0.6	...	...	...	...
IRAS 08572+3915	4.1	3.8	0.9	1.0	0.24	0.26
IRAS 09039+0503	2.6	2.2	0.9	1.0	0.41	0.45
IRAS 09116+0334	1.5	1.1	...	...	...	...
IRAS 09539+0857	3.8	3.5	1.3	1.6	0.37	0.46
IRAS 10378+1108	2.6	2.4	0.45	0.55	0.19	0.23
IRAS 10485–1447	3.0	3.0	0.8	0.95	0.27	0.32
IRAS 10494+4424	2.2	1.7	...	...	...	...
IRAS 11095–0238	3.5	3.4	0.8	0.9	0.23	0.26
IRAS 11130–2659	3.5	3.1	1.0	1.2	0.32	0.39
IRAS 12112+0305	2.3	1.7	...	...	...	...
IRAS 12127–1412	3.0	2.5	0.6	0.7	0.24	0.28
IRAS 12359–0725	2.2	1.8	0.5	0.55	0.28	0.31
IRAS 13335–2612	1.6	1.4	...	...	...	...
IRAS 14252–1550	1.6	1.4	...	...	...	...
IRAS 14348–1447	2.8	2.0	...	...	...	...
Arp 220	3.7	3.2	0.95	1.1	0.30	0.34
IRAS 16090–0139	3.0	2.6	0.7	0.8	0.27	0.31
IRAS 16468+5200	3.0	2.5	0.65	0.8	0.26	0.32
IRAS 16487+5447	2.4	1.8	...	...	...	...
IRAS 17028+5817	2.0	1.5	...	...	...	...
IRAS 17044+6720	1.7	1.8	0.4	0.5	0.22	0.28
IRAS 21329–2346	2.7	2.2	...	...	...	...
IRAS 23234+0946	2.2	1.8	...	...	...	...
IRAS 23327+2913	1.5	1.3	...	...	...	...
IRAS 00091–0738	3.2	3.2	1.05	1.25	0.33	0.39
IRAS 00456–2904	1.5	1.2	...	...	...	...
IRAS 01004–2237	0.4	0.4	...	...	...	...
IRAS 01166–0844	3.2	3.0	0.75	0.9	0.25	0.30
IRAS 01298–0744	4.3	4.0	1.1	1.35	0.27	0.34
IRAS 01569–2939	3.2	2.8	1.0	1.2	0.36	0.43
IRAS 02411+0353	1.3	1.0	...	...	...	...
IRAS 10190+1322 ^a	1.7	1.3	...	...	...	...
IRAS 11387+4116	1.6	1.1	...	...	...	...
IRAS 11506+1331	2.6	2.3	0.7	0.85	0.30	0.37
IRAS 13509+0442	2.0	1.6	...	...	...	...
IRAS 13539+2920	1.8	1.6	...	...	...	...
IRAS 14060+2919	1.5	1.0	...	...	...	...
IRAS 15206+3342	0.9	0.6	...	...	...	...
IRAS 15225+2350	2.5	2.2	0.65	0.75	0.29	0.34
IRAS 16474+3430	1.8	1.5	...	...	...	...
IRAS 20414–1651	2.4	1.9	...	...	...	...
IRAS 21208–0519	1.5	1.4	...	...	...	...
IRAS 22206–2715	2.4	1.7	...	...	...	...
IRAS 22491–1808	2.0	1.5	...	...	...	...

NOTES.—Col. (1): Object name. Col. (2): $\tau_{9.7}$ is an optical depth of the 9.7 μm silicate dust absorption feature, against a power-law continuum determined from data points at $\lambda_{\text{rest}} = 5.6$ and 7.1 μm and $\lambda_{\text{obs}} = 34\text{--}35$ μm , shown as dashed lines in Figs. 3 and 4. Col. (3): $\tau'_{9.7}$ is an optical depth of the 9.7 μm silicate dust absorption feature, against a power-law continuum determined from data points at $\lambda_{\text{rest}} = 7.1$ and 14.2 μm , shown as dotted lines in Figs. 3 and 4. Once the continuum levels are fixed, the uncertainty of $\tau'_{9.7}$ is $<5\%$ for ULIRGs with large $\tau'_{9.7}$ values (>2) and can be $\sim 10\%$ for ULIRGs with small $\tau'_{9.7}$. Col. (4): τ'_{18} is an optical depth of the 18 μm silicate dust absorption feature, against a power-law continuum determined from data points at $\lambda_{\text{rest}} = 14.2$ and 24 μm , shown as dotted lines in Figs. 3 and 4. Once the continuum is fixed, the statistic uncertainty of τ'_{18} is $<10\%$, because the value is estimated only for ULIRGs with clearly detectable 18 μm silicate absorption features. Col. (5): τ''_{18} is an optical depth of the 18 μm silicate dust absorption feature, against a power-law continuum determined from data points at $\lambda_{\text{rest}} = 14.2$ and 29 μm , shown as dash-dotted lines in Figs. 3 and 4. Once the continuum is fixed, the statistic uncertainty is $<10\%$. Col. (6): $\tau'_{18}/\tau'_{9.7}$ ratio, only for ULIRGs with clearly detectable 18 μm silicate absorption. The uncertainty is $<10\%$ once the continuum is fixed. Col. (7): $\tau''_{18}/\tau'_{9.7}$ ratio, only for ULIRGs with clearly detectable 18 μm silicate absorption.

^a To estimate the optical depths of silicate dust absorption features, SL spectra of the eastern and western nuclei are summed.

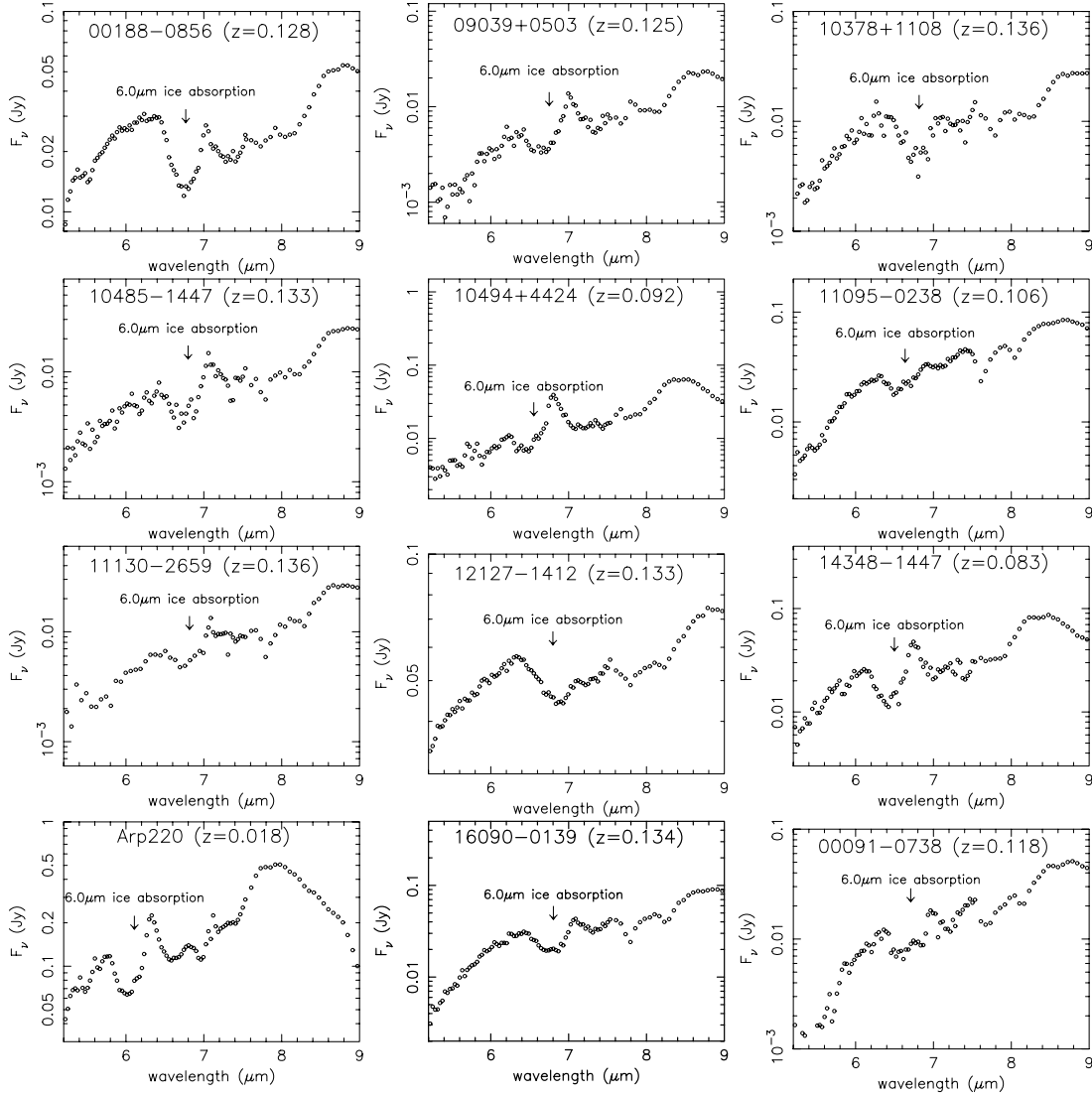


FIG. 7.—IRS spectra at $\lambda_{\text{obs}} = 5.2\text{--}9\text{ }\mu\text{m}$ for ULIRGs displaying clear $6.0\text{ }\mu\text{m}$ H_2O absorption feature (marked with “ $6.0\text{ }\mu\text{m}$ ice absorption” in the first 15 plots). Spectra of the next five ULIRGs [marked with “ $6.0\text{ }\mu\text{m}$ ice absorption (?)”] show possible signatures of the $6.0\text{ }\mu\text{m}$ ice absorption feature. Spectra of the last two ULIRGs, IRAS 17044+6720 and 00456–2904, marked with “no $6.0\text{ }\mu\text{m}$ ice absorption” are shown as examples with undetected ice absorption features. The abscissa is observed wavelength in μm in linear scale, and the ordinate is flux in Jy in decimal logarithmic scale.

For ULIRGs with large contributions from modestly obscured PAH-emitting starbursts to observed spectra, the $18\text{ }\mu\text{m}$ silicate dust absorption features are diluted by the PAH complex at $\lambda_{\text{rest}} = 16.4$ and $17.12\text{ }\mu\text{m}$, as well as H_2 emission at $\lambda_{\text{rest}} = 17.04$ and $[\text{S III}]$ emission at $\lambda_{\text{rest}} = 18.71\text{ }\mu\text{m}$ (Smith et al. 2004). Subtraction of a template spectrum of the PAH-emitting starburst introduces an additional ambiguity in the estimate of the intrinsic τ'_{18} (τ''_{18}) value toward the ULIRG’s core. To minimize the uncertainties, we investigate the $\tau'_{18}/\tau'_{9.7}$ ($\tau''_{18}/\tau'_{9.7}$) ratios only for ULIRGs where the contributions from PAH-emitting normal starbursts are small and the $18\text{ }\mu\text{m}$ silicate features are clearly seen in absorption in the observed spectra. The main aim of applying this $\tau'_{18}/\tau'_{9.7}$ ($\tau''_{18}/\tau'_{9.7}$) method to these ULIRGs is to further strengthen the case for buried AGNs previously suggested by the large $\tau'_{9.7}$ values, rather than to discover new buried AGN candidates.

To investigate the presence of a strong dust temperature gradient in a ULIRG, based on a small observed $\tau'_{18}/\tau'_{9.7}$ ($\tau''_{18}/\tau'_{9.7}$) ratio, we first have to know its intrinsic ratio in the absence of a temperature gradient. Figure 6a presents the observed profiles

of two Galactic obscured stars, GCS3-I and WR 98a (Chiar & Tielens 2006). They are thought to be mostly obscured by foreground *interstellar* medium far away from the central illuminating stars, rather than circumstellar material, so no significant temperature gradient is expected. To conservatively argue that a strong dust temperature gradient is found, based on a small observed $\tau'_{18}/\tau'_{9.7}$ ($\tau''_{18}/\tau'_{9.7}$) ratio, we have adopted the profile showing the smallest intrinsic $\tau'_{18}/\tau'_{9.7}$ ($\tau''_{18}/\tau'_{9.7}$) ratio as a template. The intrinsic ratio of GCS3-I is estimated to be $\tau'_{18}/\tau'_{9.7} = 0.3$ and $\tau''_{18}/\tau'_{9.7} = 0.35$. The ratios of WR 98a, and other Galactic obscured stars for which silicate profiles are investigated in detail (OH-IR127.8+0.0, WR 118, and WR 112; Chiar & Tielens 2006), are larger than that of GCS3-I, so that adopting the observed GCS3-I profile as a template is appropriate here. If an observed $\tau'_{18}/\tau'_{9.7}$ or $\tau''_{18}/\tau'_{9.7}$ ratio is substantially smaller than the adopted intrinsic value ($\tau'_{18}/\tau'_{9.7} = 0.3$ and $\tau''_{18}/\tau'_{9.7} = 0.35$), then the most reasonable explanation is the presence of a strong dust temperature gradient.

Table 4 (cols. [6] and [7]) summarizes the $\tau'_{18}/\tau'_{9.7}$ and $\tau''_{18}/\tau'_{9.7}$ ratios for ULIRGs showing clear $18\text{ }\mu\text{m}$ silicate absorption

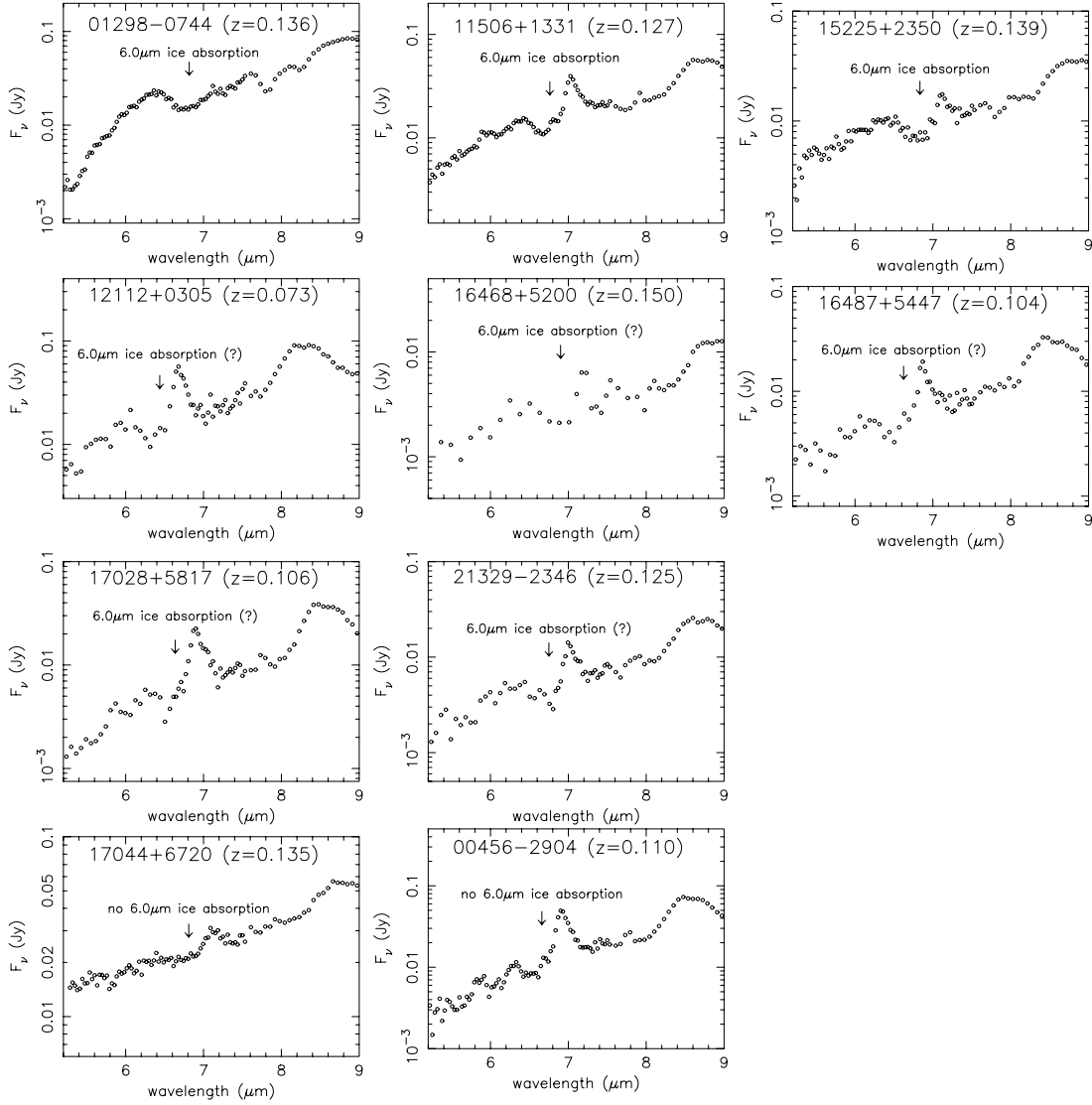


FIG. 7—Continued

(mostly $\tau'_{9.7} > 2$). For a significant fraction of such ULIRGs, the observed ratios are smaller than the conservatively assumed intrinsic one, suggesting the presence of strong dust temperature gradients. This provides additional evidence for these ULIRGs for centrally concentrated buried AGNs, previously suggested by low PAH equivalent widths and/or large $\tau'_{9.7}$ values (>2).

ULIRGs showing detectable signatures of strong dust temperature gradients are marked as open circles in Table 6 (col. [6]). The fraction of ULIRGs with detectable signatures of strong dust temperature gradients is 7/48 (15%; LINER + H II region), 6/28 (21%; LINER), and 1/20 (5%; H II region). The fraction of ULIRGs with detectable buried AGN signatures is again higher in LINER ULIRGs than H II region ULIRGs.

Again, there are several caveats. First, we adopted the observed profile of GCS3-I as the template, since the $\tau'_{18}/\tau'_{9.7}$ ratio in GCS3-I is the lowest among five Galactic sources so far investigated in detail. A source with an even smaller intrinsic $\tau'_{18}/\tau'_{9.7}$ ratio may be discovered among Galactic sources in future observations. If ULIRGs showed intrinsically smaller $\tau'_{18}/\tau'_{9.7}$ ratios than the adopted ratio from GCS3-I, then the small observed ratios in ULIRGs would not necessarily require the strong dust temperature gradients. However, for IRAS 08572+3915, a bright ULIRG

with a small $\tau'_{18}/\tau'_{9.7}$ ratio (Table 4), Spoon et al. (2006) independently found the signature of a strong dust temperature gradient, based on the fact that the observed optical depths of crystalline silicate dust absorption at $\lambda_{\text{rest}} \sim 11, 16, 18.5, 23$, and $28 \mu\text{m}$ are systematically weaker at the longer wavelengths than expected (their Fig. 2, *bottom*). A similar trend was found also in the average spectrum of ULIRGs with $\tau'_{9.7} > 2$ (Spoon et al. 2006, their Fig. 5). We thus believe that the small $\tau'_{18}/\tau'_{9.7}$ ratios observed in several ULIRGs are most naturally explained by strong dust temperature gradients.

Next, although we investigated the $\tau'_{18}/\tau'_{9.7}$ ratios only for ULIRGs with small PAH contributions from modestly obscured starbursts, these small contributions could reduce both τ'_{18} and $\tau'_{9.7}$ (§ 5.2.2), and slightly alter the $\tau'_{18}/\tau'_{9.7}$ ratios, compared to a pure spectrum of the highly obscured energy sources in the cores. However, the average spectrum of starburst galaxies shows the 30–5 μm (rest-frame) flux ratio to be ~ 900 (Brandl et al. 2006), while the majority of the ULIRGs with small $\tau'_{18}/\tau'_{9.7}$ show flux ratios larger than >900 (Figs. 1 and 2). That is, the average starburst spectral energy distribution is bluer at $\lambda_{\text{rest}} = 5\text{--}30 \mu\text{m}$ than those ULIRGs with small $\tau'_{18}/\tau'_{9.7}$. Assuming that the average spectrum is representative of the spectral energy distribution of

TABLE 5
ICE ABSORPTION

Object (1)	6.0 μm (2)	$\tau_{6.0}$ (3)	3.1 μm (4)	$\tau_{3.1}$ (5)
IRAS 00188–0856.....	○	0.9	○	1.8
IRAS 00482–2721.....	X	...	...	...
IRAS 03250+1606.....	X	...	○	0.8
IRAS 04103–2838.....	X	...	...	...
IRAS 08572+3915.....	X	...	X	...
IRAS 09039+0503.....	○	0.6	○	0.8
IRAS 09116+0334.....	X	...	○	0.4
IRAS 09539+0857.....	X	...	X	...
IRAS 10378+1108.....	○	0.8	○	0.6
IRAS 10485–1447.....	○	0.5	○	0.8
IRAS 10494+4424.....	○	0.4	○	1.0
IRAS 11095–0238.....	○	0.5	X	...
IRAS 11130–2659.....	○	0.5	...	...
IRAS 12112+0305.....	△	...	X	...
IRAS 12127–1412.....	○	0.4	○	0.4
IRAS 12359–0725.....	X	...	○	0.5
IRAS 13335–2612.....	X	...	...	...
IRAS 14252–1550.....	X	...	○	0.7
IRAS 14348–1447.....	○	0.9	○	0.5
Arp 220.....	○	0.6	...	...
IRAS 16090–0139.....	○	0.5	○	0.8
IRAS 16468+5200.....	△	...	X	...
IRAS 16487+5447.....	△	...	○	1.2
IRAS 17028+5817.....	△	...	○	0.6
IRAS 17044+6720.....	X	...	X	...
IRAS 21329–2346.....	△	...	○	1.0
IRAS 23234+0946.....	X	...	X	...
IRAS 23327+2913.....	X	...	X	...
IRAS 00091–0738.....	○	0.5	...	...
IRAS 00456–2904.....	X	...	...	...
IRAS 01004–2237.....	X	...	...	...
IRAS 01166–0844.....	X	...	...	...
IRAS 01298–0744.....	○	0.5	...	...
IRAS 01569–2939.....	X	...	...	...
IRAS 02411+0353.....	X	...	...	...
IRAS 10190+1322E.....	X	...	X	...
IRAS 10190+1322W.....	X	...	X	...
IRAS 11387+4116.....	X	...	X	...
IRAS 11506+1331.....	○	0.5	○	1.0
IRAS 13509+0442.....	X	...	X	...
IRAS 13539+2920.....	X	...	X	...
IRAS 14060+2919.....	X	...	X	...
IRAS 15206+3342.....	X	...	X	...
IRAS 15225+2350.....	○	0.5	○	0.4
IRAS 16474+3430.....	X	...	○	0.8
IRAS 20414–1651.....	X	...	X	...
IRAS 21208–0519.....	X	...	X	...
IRAS 22206–2715.....	X	...	...	...
IRAS 22491–1808.....	X	...	...	...

NOTES.—Col. (1): Object name. Col. (2): Detection or nondetection of the 6.0 μm ice absorption feature: (○) clearly detected; (△) possibly detected; (X) nondetected. The nondetection includes two cases: (1) the absorption feature is indeed weak or absent in high signal-to-noise ratios spectra or (2) the presence of ice absorption is not clear due to relatively large scatter in spectra or to the difficulty in separating the ice absorption and nearby PAH emission features. Col. (3): Optical depth of the 6.0 μm ice absorption feature. Ellipses signify the depth is not measurable. Since data scatter in spectra is not the same for individual ULIRGs, there is no fixed detection threshold. The uncertainty is ~ 0.1 for detected sources. Col. (4): Detection or nondetection of the 3.1 μm ice absorption feature in the 2.8–4.1 μm spectra of Imanishi et al. (2006a): (○) clearly detected; (X) nondetected. Ellipses indicate no data. Col. (5): Optical depth of the 3.1 μm ice absorption feature. Ellipses signify the depth is not measurable because (1) it is weak, (2) data scatter is large in spectra, and (3) 2.8–4.1 μm spectra are lacking. Again there is no fixed detection threshold.

the modestly obscured starbursts in these ULIRGs, this possible effect increases (not decreases) the observed $\tau'_{18}/\tau'_{9.7}$ ratios, compared to an uncontaminated spectrum of the highly obscured core. Thus, the observed small $\tau'_{18}/\tau'_{9.7}$ ratios in these ULIRGs are still best explained by strong dust temperature gradients around buried AGNs, with centrally concentrated energy source geometry.

Finally, of course, silicate dust grains in the cores of ULIRGs may have intrinsically larger $\tau'_{18}/\tau'_{9.7}$ ratios than the conservatively assumed one. In this case, if a ULIRG shows a decreased $\tau'_{18}/\tau'_{9.7}$ ratio compared to the intrinsic one, due to a dust temperature gradient, but still has an *observed* ratio similar to the assumed one, then such a ULIRG would not be classified as a buried AGN.

5.2.4. Comparison of the Strength of Silicate Absorption and PAH Emission Features

Figure 10a compares the optical depth of the 9.7 μm silicate dust absorption ($\tau'_{9.7}$) and the rest-frame equivalent widths of the 6.2 μm PAH emission feature ($\text{EW}_{6.2\text{PAH}}$). With the exception of the two outliers with both small $\text{EW}_{6.2\text{PAH}}$ and small $\tau'_{9.7}$ (IRAS 01004–2237 and 17044+6720, both weakly obscured AGNs), there is a weak anticorrelation, and ULIRGs with $\tau'_{9.7} > 2$ (i.e., buried AGN candidates) tend to show small $\text{EW}_{6.2\text{PAH}}$ compared to starburst-like ULIRGs with small $\tau'_{9.7}$ (< 2).

A highly obscured ($A_V \gg 30$ mag) normal starburst galaxy with mixed dust/source geometry (scenario A) predicts a relatively flat distribution (i.e., relatively constant $\text{EW}_{6.2\text{PAH}}$ values with varying $\tau'_{9.7}$), because the PAH equivalent width is unchanged by dust extinction (§ 1). On the other hand, scenario B, where (1) the PAH emission comes from modestly obscured ($A_V < 20$ –30 mag) starbursts exterior to the deeply obscured ULIRG cores and (2) the large $\tau'_{9.7}$ values reflect dust column density toward the buried AGNs in the cores, predicts an anticorrelation between $\tau'_{9.7}$ and $\text{EW}_{6.2\text{PAH}}$. As the contribution from the AGN-produced PAH-free continuum to the observed flux increases, the $\text{EW}_{6.2\text{PAH}}$ value decreases. At the same time, the increasing contribution from the buried AGN emission (with centrally concentrated energy source geometry) results in the increase of the observed $\tau'_{9.7}$ value, compared to a pure starburst with mixed source/dust geometry. The trend in Figure 10a (see also Spoon et al. 2007) supports the scenario B for ULIRGs with $\tau'_{9.7} > 2$.

Figure 10b compares $\tau'_{9.7}$ and the rest-frame equivalent widths of the 11.3 μm PAH emission feature ($\text{EW}_{11.3\text{PAH}}$). Compared to Figure 10a, the decreasing trend of $\text{EW}_{11.3\text{PAH}}$ at $\tau'_{9.7} > 2$ appears slight at best. In scenario B, the flux depression of the AGN-produced continuum flux at $\lambda_{\text{rest}} = 11.3$ μm is particularly large due to the strong 9.7 μm silicate absorption. However, the 11.3 μm PAH emission flux is less affected by this absorption because the PAH emission comes from starburst regions outside the core. Consequently, the decrease of the $\text{EW}_{11.3\text{PAH}}$ value with increasing $\tau'_{9.7}$ value is small compared to $\text{EW}_{6.2\text{PAH}}$, explaining the trends in both Figs. 10a and 10b.

Figure 11 compares the $\tau'_{9.7}$ values with the 11.3–6.2 μm PAH luminosity ratio. In scenario A, the 11.3–6.2 μm flux ratio should decrease with increasing $\tau'_{9.7}$, as actually seen in normal starburst galaxies (Brandl et al. 2006), because the flux depression of the 11.3 μm PAH emission inside the strong 9.7 μm absorption is larger. On the other hand, in scenario B, no correlation is expected because both the 6.2 and 11.3 μm PAH emission come from the modestly obscured starbursts, unrelated to the observed $\tau'_{9.7}$ values. Figure 11 shows no significant correlation, supporting scenario B.

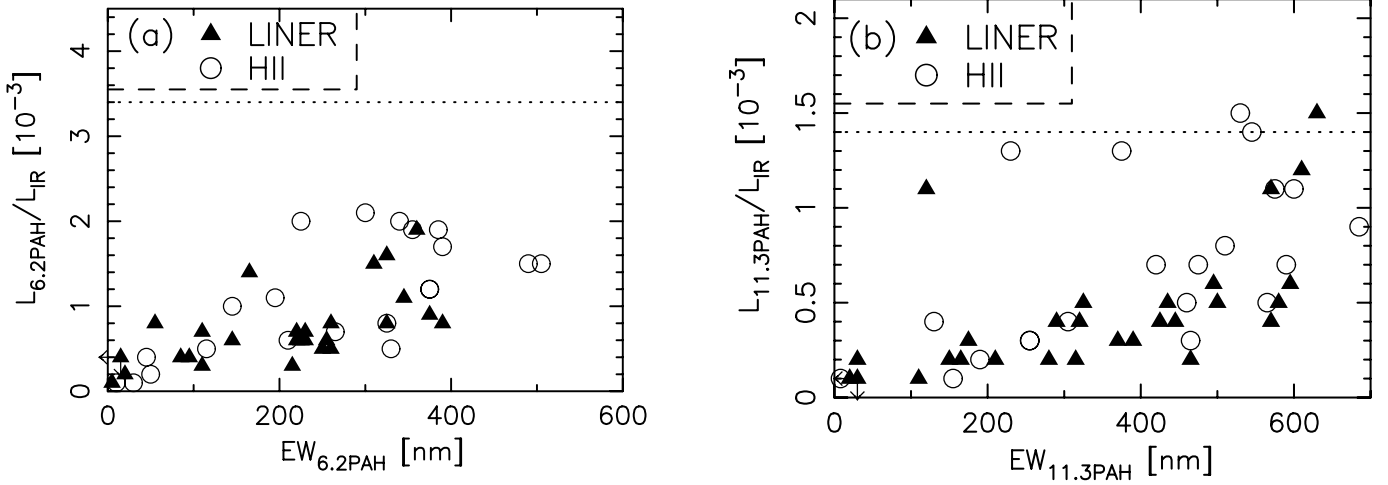


FIG. 8.—(a) Comparison of the rest-frame equivalent width of the 6.2 μm PAH emission (*abscissa*) and 6.2 μm PAH-to-infrared luminosity ratio (*ordinate*). *Filled triangle*: ULIRGs classified optically as LINERs. *Open circle*: ULIRGs classified optically as H II regions. The horizontal dotted line indicates the ratio for normal starburst galaxies with modest dust obscuration ($L_{6.2\text{PAH}}/L_{\text{IR}} = 3.4 \times 10^{-3}$; see § 5.1). In the upper left small area separated by the dashed line, there are no data points. (b) The *abscissa* is the rest-frame equivalent width of the 11.3 μm PAH emission, and the *ordinate* is the 11.3 μm PAH-to-infrared luminosity ratio. The horizontal dotted line indicates the ratio for normal starburst galaxies with modest dust obscuration ($L_{11.3\text{PAH}}/L_{\text{IR}} = 1.4 \times 10^{-3}$; see § 5.1). The larger scatter of $L_{11.3\text{PAH}}/L_{\text{IR}}$ in this figure than for $L_{6.2\text{PAH}}/L_{\text{IR}}$ in (a) may be due to intrinsically larger scatter of the 11.3 μm PAH strengths in normal starburst galaxies (see § 5.2.1).

5.2.5. Is the Buried AGN Scenario Unique?

ULIRGs with large $\tau'_{9.7}$ (>2) and small $\tau'_{18}/\tau'_{9.7}$ ratios are naturally explained by the presence of buried AGNs (scenario B). We here briefly comment on two alternative scenarios that are possible but we believe are unlikely.

First, in a normal starburst with mixed dust/source geometry, if each star is deeply buried in a large amount of *circumstellar* dust cocoon, such as a protostar, then a strong dust absorption feature ($\tau'_{9.7} > 2$) could be produced. However, the timescale of such a protostar phase is only $<10^6$ yr (Lada 1987), which is much less than the typical timescale of a normal starburst ($\sim 10^7$ – 10^8 yr; Moorwood 1996) or even the maximum starburst ($\sim 10^7$ – 10^8 yr; Elmegreen 1999), or a ULIRG timescale ($\sim 10^8$ – 10^9 yr; Murphy et al. 1996). It is therefore very unlikely that all dust-heating stars are protostars in a significant fraction of ULIRGs. Furthermore, this scenario produces no strong dust temperature gradient and so cannot reduce the observed $\tau'_{18}/\tau'_{9.7}$ ratio.

Second, it may be argued that exceptionally centrally concentrated starbursts remain a possibility (Fig. 1e). However, while starburst galaxies with $L_{\text{IR}} = 10^{9.7}$ – $10^{11.6} L_{\odot}$ are well explained by the mixed dust/source geometry (Brandl et al. 2006; see § 5.2.4 of our paper), there is no reason why starbursts at the cores of ULIRGs ($L_{\text{IR}} > 10^{12} L_{\odot}$) suddenly change their properties to exceptionally centrally concentrated nature (Fig. 1e). Furthermore, Soifer et al. (2000) studied the cores of seven nearby ULIRGs and found that the observed emission surface brightnesses in the majority (six out of seven) of these ULIRGs' cores are very high ($\sim 10^{13} L_{\odot} \text{ kpc}^{-2}$). Even if the energy sources are *uniformly* distributed over the cores (<500 pc in size), the observed surface brightnesses are close to the maximum values seen in the cores of Galactic H II regions (Soifer et al. 2000). If an *exceptionally centrally concentrated* starburst were the luminosity source in the core of a ULIRG, the emission surface brightness of the starburst would have to be extremely high, $>10^{14} L_{\odot} \text{ kpc}^{-2}$. Among known star-forming regions, such a high surface brightness could only barely be achieved with a super star cluster, where many young stars are formed within an area of less than a few parsec (Gorjian et al. 2001). The *absolute luminosity* of such super star clusters is, however, only of order 10^9 – $10^{10} L_{\odot}$, and so the

super star cluster scenario may work for the centrally concentrated energy sources of low-luminosity ($L_{\text{IR}} \sim 10^{10} L_{\odot}$) galaxy nuclei (Imanishi & Nakanishi 2006). For the centrally concentrated energy sources at ULIRG cores, both *high emission surface brightness* and *large absolute luminosity* ($\sim 10^{12} L_{\odot}$) are required, making the buried AGN scenario the most plausible choice (Soifer et al. 2003, their § 6). In fact, known AGN populations (i.e., quasars) meet these requirements without difficulty, while for the starburst scenario, of order 100–1000 super star clusters must be concentrated in a very compact region and such a phenomenon has not been clearly confirmed from observations. We thus believe that buried AGNs are the most natural explanation for the centrally concentrated energy sources at ULIRGs' cores. The presence of luminous buried AGNs in many nearby ULIRGs agrees with the suggestion by Takeuchi et al. (2003) that the observed lump of the local infrared 60 μm luminosity function at the highest luminosity end is due to the AGN contribution.

5.2.6. Combination of Energy Diagnostic Methods

Table 6 (col. [7]) summarizes the strengths of the detected buried AGN signatures in *Spitzer* IRS 5–35 μm spectra based on the three methods: (1) low PAH equivalent width, (2) large $\tau'_{9.7}$ value, and (3) small $\tau'_{18}/\tau'_{9.7}$ and $\tau'_{18}/\tau'_{9.7}$ ratio. When buried AGN signatures in individual ULIRGs are consistently found using all or most of the methods, the ULIRGs are classified as *very strong* buried AGN candidates, marked with open double circles. When buried AGNs signatures are seen in the first method (particularly based on the observed $\text{EW}_{6.2\text{PAH}}$ value) and/or in the third method, then the ULIRGs are classified as *strong* AGN candidates (*open circles*). When the signatures are detected only in the second method and/or low observed $\text{EW}_{11.3\text{PAH}}$ or corrected $\text{EW}_{6.2\text{PAH}}$ values (first method), the ULIRGs are classified as *possible* buried AGN candidates (*open triangles*), because (1) a normal starburst nucleus obscured by foreground dust in the host galaxy (Fig. 1d) cannot be ruled out completely and (2) the 11.3 μm PAH emission strengths in normal starbursts have intrinsically larger scatter (§ 5.2.1). The fraction of ULIRGs with *strong* buried AGN signatures is 16/48 (33%; LINER + H II region), 10/28 (36%; LINER), and 6/20 (30%; H II regions). When ULIRGs with

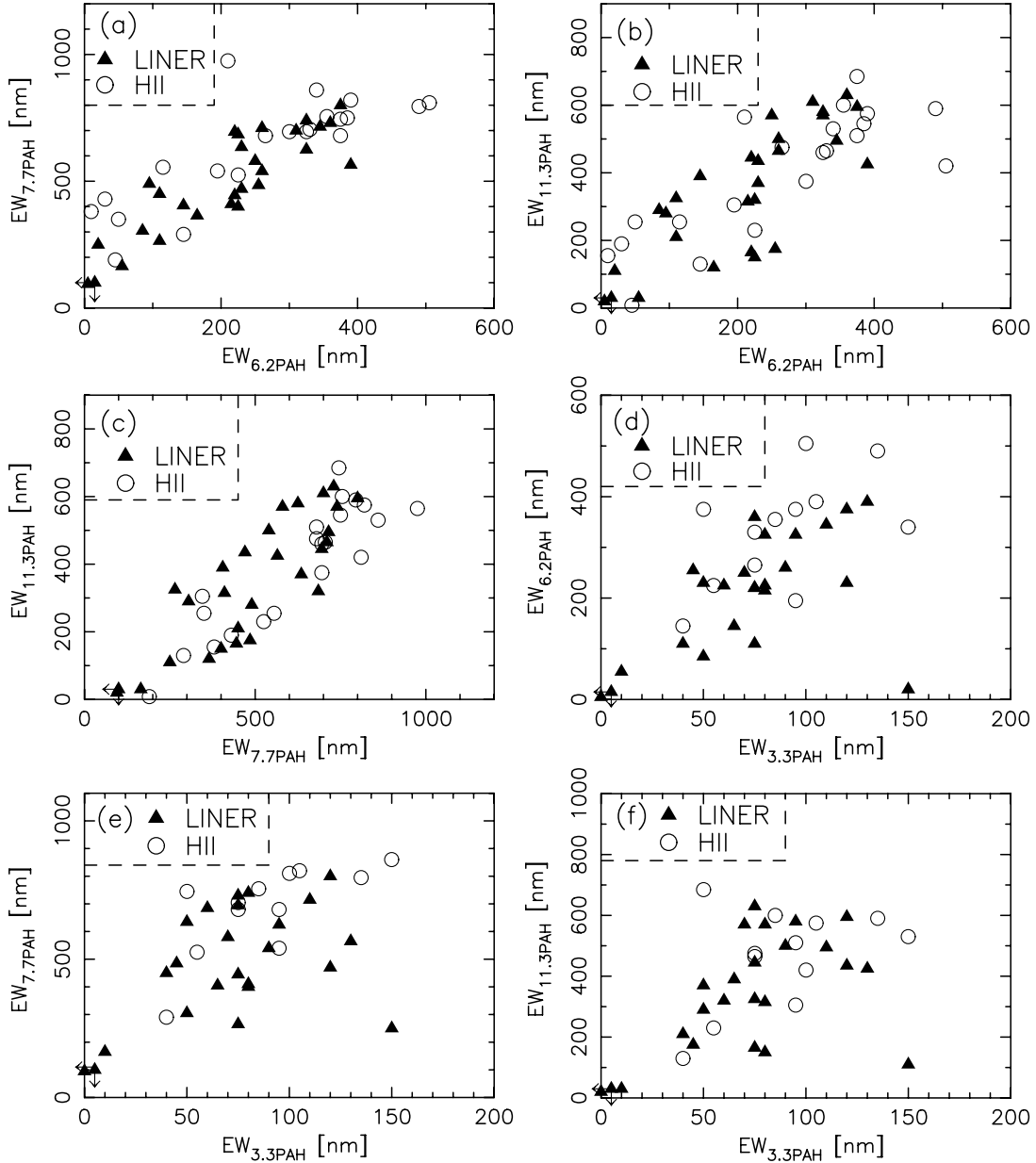


FIG. 9.—(a) Comparison of the rest-frame equivalent widths of PAH emission features in nm. The abscissa is the $6.2\ \mu\text{m}$ PAH, and the ordinate is the $7.7\ \mu\text{m}$ PAH emission. *Filled triangle*: ULIRGs classified optically as LINERs. *Open circle*: ULIRGs classified optically as H II regions. In the upper left small area separated with the dashed line, there are no data points. (b) The abscissa is the $6.2\ \mu\text{m}$ PAH, and the ordinate is the $11.3\ \mu\text{m}$ PAH emission. (c) The abscissa is the $7.7\ \mu\text{m}$ PAH, and the ordinate is the $11.3\ \mu\text{m}$ PAH emission. (d) The abscissa is the $3.3\ \mu\text{m}$ PAH, and the ordinate is the $6.2\ \mu\text{m}$ PAH emission. (e) The abscissa is the $3.3\ \mu\text{m}$ PAH, and the ordinate is the $7.7\ \mu\text{m}$ PAH emission. (f) The abscissa is the $3.3\ \mu\text{m}$ PAH, and the ordinate is the $11.3\ \mu\text{m}$ PAH emission. In the last three plots, the average $\text{EW}_{3.3\text{PAH}}$ value of both nuclei is plotted for IRAS 12112+0305 and 14348–1447 (see Table 3).

possible buried AGN signatures are included, the fraction increases to 26/48 (54%; LINER + H II region), 19/28 (68%; LINER), and 7/20 (35%; H II region). Our careful look at the infrared 5–35 μm spectra suggests that luminous buried AGNs reside in a much higher fraction of optical non-Seyfert ULIRGs than previous estimates (Taniguchi et al. 1999; Lutz et al. 1999).

We stress that the above fraction of luminous buried AGNs is only for ULIRGs classified optically as LINERs and H II regions (i.e., non-Seyferts), which comprise 70% of ULIRGs. Given that 30% of ULIRGs show optical Seyfert signatures, indicative of luminous AGNs surrounded by dusty tori, and that the AGNs are estimated to be energetically dominant in these optical Seyfert ULIRGs (Veilleux et al. 1999b), the total fraction of ULIRGs possessing luminous AGNs is at least 50%, and possibly >65%,

in the whole ULIRG sample at $z < 0.15$, including both optical Seyfert and non-Seyfert objects.

Finally, in all of the methods employed by us (§ 5.2.1–5.2.3 and 5.2.6), ULIRGs classified optically as LINERs tend to display signatures of luminous buried AGNs more frequently than those classified optically as H II regions. For starburst/buried AGN composite ULIRGs (Fig. 1c), it is likely that the optical LINER or H II region classifications are largely affected by the properties of the modestly obscured starbursts at the exteriors of the ULIRG cores, rather than buried AGN-related emission, because optical observations can probe only the surfaces of dusty objects. If shocks or superwinds are important in the modestly obscured starbursts, the optical classification will be LINERs. If emission from H II regions is dominant, the ULIRGs will be

TABLE 6
BURIED AGN SIGNATURES

OBJECT (1)	Spitzer IRS 5–35 μm						3–4 μm (8)
	EW _{6.2PAH} (2)	EW _{7.7PAH} (3)	EW _{11.3PAH} (4)	$\tau'_{9.7}$ (5)	T-Gradient (6)	Total (7)	
IRAS 00188–0856	○	X	X	○	○	⊙	○
IRAS 00482–2721	X	X	X	○	X	△	...
IRAS 03250+1606	X	X	X	X	X	X	○
IRAS 04103–2838	○	X	○	X	X	○	...
IRAS 08572+3915	○	○	○	○	○	⊙	⊙
IRAS 09039+0503	X	X	X	○	X	△	○
IRAS 09116+0334	X	X	X	X	X	X	○
IRAS 09539+0857	○	X	X	○	X	○	X
IRAS 10378+1108	○	X	X	○	○	○	○
IRAS 10485–1447	X (○)	X	○	○	△	△	○
IRAS 10494+4424	X	X	X	X	X	X	○
IRAS 11095–0238	○	X	○	○	○	⊙	X
IRAS 11130–2659	○	X	X	○	X	○	X
IRAS 12112+0305	X	X	X	X	X	X	X
IRAS 12127–1412	○	○	○	○	○	⊙	⊙
IRAS 12359–0725	X	X	○	X	△	△	○
IRAS 13335–2612	X	X	X	X	X	X	...
IRAS 14252–1550	X	X	X	X	X	X	○
IRAS 14348–1447	X (○)	X	X	X	X	△	○
Arp 220	X (○)	X	X	○	X	△	X
IRAS 16090–0139	○	X	X	○	△	○	○
IRAS 16468+5200	X	X	X	○	△	△	X
IRAS 16487+5447	X	X	X	X	X	X	○
IRAS 17028+5817	X	X	X	X	X	X	○
IRAS 17044+6720	○	○	○	X (...) ^a	○	⊙	⊙
IRAS 21329–2346	X	X	X	○	X	△	○
IRAS 23234+0946	X	X	X	X	X	X	X
IRAS 23327+2913	X	X	○	X	X	△	X
IRAS 00091–0738	○	X	X	○	X	○	...
IRAS 00456–2904	X	X	X	X	X	X	...
IRAS 01004–2237	○	○	○	X (...) ^a	X	⊙	...
IRAS 01166–0844	○	X	○	○	○	⊙	...
IRAS 01298–0744	○	X	○	○	X	○	...
IRAS 01569–2939	○	X	X	○	X	○	...
IRAS 02411+0353	X	X	X	X	X	X	...
IRAS 10190+1322	X ^b	X ^b	X ^b	X	X	X	X
IRAS 11387+4116	X	X	X	X	X	X	X
IRAS 11506+1331	X (○)	X	X	○	X	△	○
IRAS 13509+0442	X	X	X	X	X	X	X
IRAS 13539+2920	X	X	X	X	X	X	X
IRAS 14060+2919	X	X	X	X	X	X	X
IRAS 15206+3342	X	X	X	X	X	X	X
IRAS 15225+2350	○	X	○	○	X	○	○
IRAS 16474+3430	X	X	X	X	X	X	○
IRAS 20414–1651	X	X	X	X	X	X	X
IRAS 21208–0519	X	X	X	X	X	X	X
IRAS 22206–2715	X	X	X	X	X	X	...
IRAS 22491–1808	X	X	X	X	X	X	...

NOTES.—Col. (1): Object name. Col. (2): Buried AGN signatures based on the low equivalent width of the 6.2 μm PAH emission ($\text{EW}_{6.2\text{PAH}} < 180 \text{ nm}$; § 5.2.1): (○) present; (X) none. When the corrected $\text{EW}_{6.2\text{PAH}}$ value becomes less than 180 nm, the symbol “○” is shown in parentheses. Col. (3): Buried AGN signatures based on the low equivalent width of the 7.7 μm PAH emission ($\text{EW}_{7.7\text{PAH}} < 230 \text{ nm}$; § 5.2.1): (○) present; (X) none; Col. (4): Buried AGN signatures based on the low equivalent width of the 11.3 μm PAH emission ($\text{EW}_{11.3\text{PAH}} < 200 \text{ nm}$; § 5.2.1): (○) present; (X) none. Col. (5): Buried AGN signatures based on the large $\tau'_{9.7}$ value (> 2 ; § 5.2.2): (○) present; (X) none. Col. (6): Buried AGN signatures based on the small $\tau'_{18}/\tau'_{9.7}$ and $\tau''_{18}/\tau'_{9.7}$ ratios: (○) present; (△) possibly present; (X) none. Considering the uncertainty of the ratio with $< 10\%$ (Table 4), ULIRGs showing $\tau'_{18}/\tau'_{9.7} < 0.27$ and $\tau''_{18}/\tau'_{9.7} < 0.32$ are taken as sources with detectable signatures of strong dust temperature gradients (§ 5.2.3). Col. (7): Buried AGN signatures from combined methods in Col. (2)–(6) based on *Spitzer* IRS spectra. ⊙: very strong; (○) strong; (△) possibly present; (X) none. Please see § 5.2.6 for more details. Col. (7): Buried AGN signatures based on infrared 3–4 μm spectroscopic diagnostic by Imanishi et al. (2006a): (⊙) very strong; (○) strong; (X) none; (...) no data.

^a The small $\tau'_{9.7}$ value is due to weak obscuration of a luminous AGN, rather than the normal starburst nature with mixed dust/source geometry.

^b PAH equivalent widths are derived for both nuclei separately. No AGN signatures are found in both nuclei.

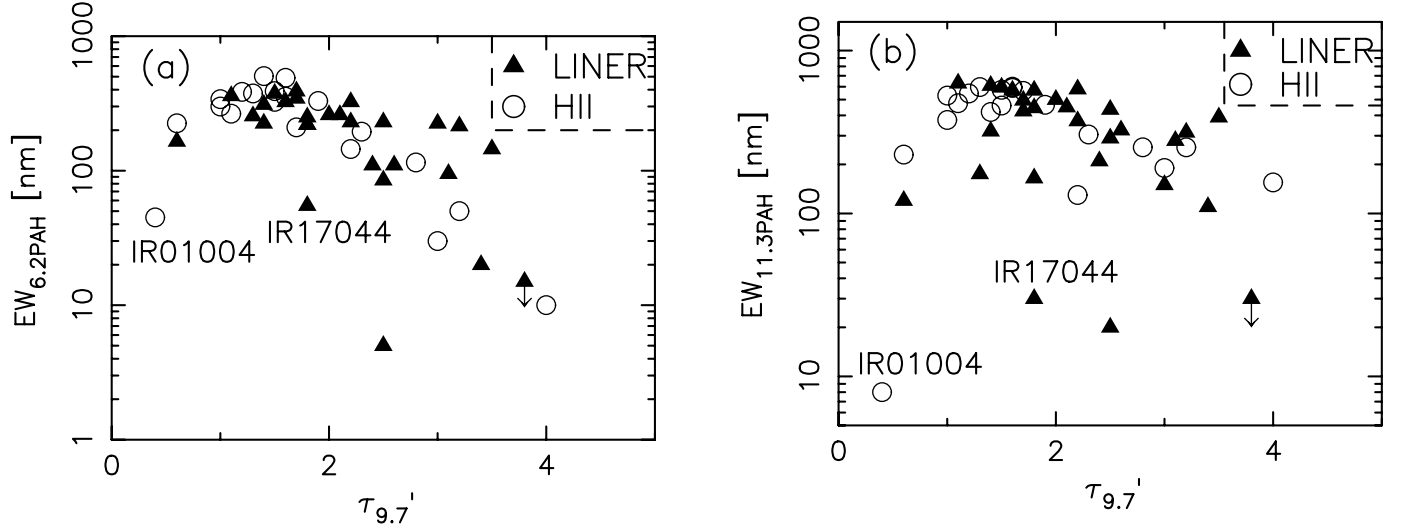


FIG. 10.—(a) Comparison of the strength of the $9.7 \mu\text{m}$ silicate dust absorption ($\tau'_{9.7}$) (abscissa) and the rest-frame equivalent widths of the $6.2 \mu\text{m}$ PAH emission feature (ordinate). Filled triangle: ULIRGs classified optically as LINERs. Open circle: ULIRGs classified optically as H II regions. In the upper right small area separated with the dashed line, there are no data points. (b) The ordinate is the rest-frame equivalent width of the $11.3 \mu\text{m}$ PAH emission feature. In both plots, the two outliers IRAS 01004–2237 and 17044+6720 (weakly obscured AGNs; see § 5.2.4) are indicated as “IR01004” and “IR17044,” respectively. For IRAS 10190+1322, the average value of both nuclei (see Table 3) are plotted for PAH equivalent widths in the ordinate.

classified as H II regions. In this case, it is not obvious why two phenomena, a luminous buried AGN at the center and shocks in the surface starburst, are correlated. Veilleux et al. (1995, 1999a) found that the emission probed in the optical is dustier in LINER ULIRGs than in H II region ULIRGs. In a dusty starburst, shock-related emission can be relatively important in the optical compared to the emission from H II regions themselves, resulting in optical LINER classification. When a luminous AGN is placed at the center of a less dusty starburst, classified optically as an H II region, the AGN emission is more easily detectable in the optical spectrum, making such an object an optical Seyfert. On the other hand, when a luminous AGN is placed at the center of a dusty starburst classified optically as a LINER, the AGN emission

is more elusive in the optical spectrum, so that such an object is classified as an optical non-Seyfert. The observed higher fraction of optically elusive *buried* AGNs in optical LINER ULIRGs compared to H II region ULIRGs can qualitatively be explained by this scenario.

5.2.7. Absorption-corrected Luminosities of Buried AGNs

For ULIRGs that show buried AGN signatures and also small PAH equivalent widths, the observed fluxes are largely ascribed to AGN-heated dust continuum emission. The resulting luminosity is conserved at each temperature (Fig. 2). We can *quantitatively* estimate the absorption-corrected intrinsic dust luminosity heated by the AGN (whole infrared), or the absorption-corrected AGN energetic radiation luminosity (X-ray–UV–optical), based on the observed fluxes at $\lambda_{\text{rest}} \sim 10 \mu\text{m}$ and the dust column density toward the $10 \mu\text{m}$ continuum-emitting regions inferred from $\tau'_{9.7}$. Since in a buried AGN, the $10 \mu\text{m}$ continuum emission regions are dominated by fixed layers (§ 5.2.3; Fig. 2), the foreground screen dust absorption model is applicable, and the dust extinction corrections are straightforward.

To do this, we assume that $\tau'_{9.7}$ and the extinction at $\lambda_{\text{rest}} = 8$ or $13 \mu\text{m}$ continuum just outside the $9.7 \mu\text{m}$ silicate feature (A_{cont}) are related with $\tau'_{9.7}/A_{\text{cont}} \sim 2.3$ (Rieke & Lebofsky 1985). Table 7 summarizes the absorption-corrected AGN luminosities for selected ULIRGs with both strong buried AGN signatures and low PAH equivalent widths. The flux attenuation of the 8 or $13 \mu\text{m}$ continuum outside the $9.7 \mu\text{m}$ silicate feature ranges from a factor of 1.2 (IRAS 01004–2237; $\tau'_{9.7} \sim 0.4$) to 5 (IRAS 01298–0744; $\tau'_{9.7} \sim 4$). The absorption-corrected AGN luminosities are $(1\text{--}5.5) \times 10^{45} \text{ ergs s}^{-1}$, which are of order $\sim 10^{12} L_{\odot}$ ($= 3.8 \times 10^{45} \text{ ergs s}^{-1}$).

The ambiguity of the derived absorption-corrected AGN luminosities come from (1) absolute flux calibration error of the *Spitzer* IRS spectra and slit loss of the compact AGN emission; (2) possibly different dust extinction curves from the Galactic one that we assume; and (3) a small contribution from PAH-emitting, modestly obscured starbursts, which decreases the $\tau'_{9.7}$ values and increases the continuum flux level, compared to pure emission from the buried AGNs. The first error is not significant,

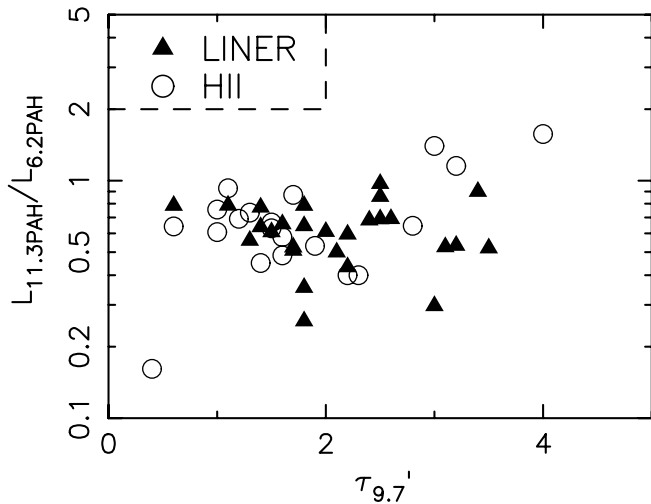


FIG. 11.—Comparison of the strength of the $9.7 \mu\text{m}$ silicate dust absorption ($\tau'_{9.7}$) (abscissa) and $11.3\text{--}6.2 \mu\text{m}$ PAH luminosity ratio (ordinate). Filled triangle: ULIRGs classified optically as LINERs. Open circle: ULIRGs classified optically as H II regions. In the upper left small area separated with the dashed line, there are no data points. For IRAS 10190+1322, the $11.3\text{--}6.2 \mu\text{m}$ PAH luminosity ratio in the ordinate is derived by combining spectra of both nuclei. IRAS 08572+3915 is not plotted, because both the 6.2 and $11.3 \mu\text{m}$ PAH emission luminosities are upper limits (see Table 3).

TABLE 7
ABSORPTION-CORRECTED AGN LUMINOSITY

Object (1)	$L(\text{AGN})$ ($10^{45} \text{ ergs s}^{-1}$) (2)
IRAS 00188–0856	1.5
IRAS 04103–2838	1.5
IRAS 08572+3915	5.5
IRAS 10378+1108	1.0
IRAS 10485–1447	1.0
IRAS 11095–0238	2.5
IRAS 11130–2659	1.0
IRAS 12127–1412	2.5
IRAS 16090–0139	2.0
IRAS 17044+6720	2.0
IRAS 00091–0738	1.5
IRAS 01004–2237	1.5
IRAS 01166–0844	1.0
IRAS 01298–0744	3.5
IRAS 15225+2350	1.5

NOTES.—Col. (1): Object name. Col. (2): Absorption-corrected intrinsic luminosity of the AGN energetic radiation (X-ray–UV–optical) in units of $10^{45} \text{ ergs s}^{-1}$.

approximately 20% or so. The second uncertainty is more difficult to quantify, but our adoption of $\tau'_{9.7}$ reduces the effects of dust grain size change (§ 5.2.2). For the third source of uncertainty, the effect of the decreased $\tau'_{9.7}$ values is more important than that of the increased continuum flux level, because the AGN-produced flux at $\lambda_{\text{rest}} \sim 9.7 \mu\text{m}$ is more highly attenuated. Assume for example that the modestly obscured starburst emission contributes 10% of the observed flux at $\lambda_{\text{rest}} = 8$ and $13 \mu\text{m}$, outside the $9.7 \mu\text{m}$ silicate feature. The buried AGN-produced continuum flux is then only 10% smaller than the observed flux. For ULIRGs with $\tau'_{9.7} > 2$ (i.e., buried AGN candidates we are now discussing), since the attenuation of the observed flux at $\lambda_{\text{rest}} = 9.7 \mu\text{m}$ is a factor of >7 compared to the nearby continuum, more than 50% of the continuum at $\lambda_{\text{rest}} = 9.7 \mu\text{m}$ is produced by the starburst emission. The true, uncontaminated $\tau'_{9.7}$ value toward the buried AGN is larger by a factor of >0.7 than the observed value, so that the correction factor of dust extinction for the AGN-heated dust continuum is larger by $>30\%$ ($\tau'_{9.7}/A_{\text{cont}} \sim 2.3$). Therefore, the small starburst contamination generally results in the underestimate of the intrinsic absorption-corrected AGN luminosity. Despite these uncertainties, the rough agreement between the absorption-corrected buried AGN luminosities and the observed infrared luminosities of these ULIRGs [$(1-3) \times 10^{12} L_{\odot}$] makes it reasonable for us to argue that the putative AGNs could quantitatively account for the bulk of the observed infrared luminosities, at least for ULIRGs with strong buried AGN signatures in *Spitzer* IRS spectra.

5.3. Comparison with Infrared 3–4 μm Spectra

5.3.1. Ice Absorption Feature

H_2O ice absorption is also present at $3.1 \mu\text{m}$ (due to a stretching mode) and has been detected in 3–4 μm spectra of many ULIRGs in this sample (Imanishi & Maloney 2003; Imanishi et al. 2006a; Risaliti et al. 2006). Comparison of the 3.1 and $6.0 \mu\text{m}$ ice absorption features can provide a better understanding of the properties of ice in ULIRGs.

Column (4) of Table 5 summarizes the detection or non-detection of the $3.1 \mu\text{m}$ ice feature. For the detections its optical depth is shown in column (5). Many ULIRGs have clear detec-

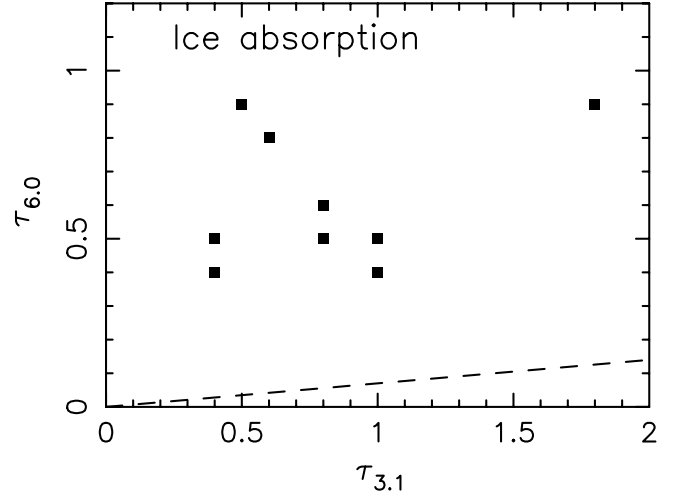


FIG. 12.—Comparison of the optical depths of the H_2O ice absorption features at $3.1 \mu\text{m}$ (abscissa) and $6.0 \mu\text{m}$ (ordinate). The dashed line is the expected ratio from a laboratory H_2O ice sample at 10 K (Keane et al. 2001).

tions of both the 3.1 and $6.0 \mu\text{m}$ ice absorption features, reinforcing the picture that the energy sources in many ULIRGs are obscured by a large amount of ice-covered dust grains.

Figure 12 compares the optical depths of the $3.1 \mu\text{m}$ ($\tau_{3.1}$) and $6.0 \mu\text{m}$ ($\tau_{6.0}$) H_2O ice absorption features, for ULIRGs clearly showing both features. The $\tau_{6.0}/\tau_{3.1}$ ratios are several times larger than expected from laboratory data. Systematically larger $\tau_{6.0}/\tau_{3.1}$ ratios compared to the laboratory prediction have previously been found in Galactic obscured stars, but the reason is unclear (Gibb et al. 2000; Keane et al. 2001). This trend may be common in astronomical obscured sources. Our poor understanding of ice absorption features hampers the search for buried AGNs with centrally concentrated energy source geometry based on small $\tau_{6.0}/\tau_{3.1}$ ratios, as was done based on the small $\tau'_{18}/\tau'_{9.7}$ ratio (§ 5.2.3).

5.3.2. PAH Emission Features

The $3.3 \mu\text{m}$ PAH emission feature assigned to the aromatic CH stretch (Allamandola et al. 1989) was detected in the majority of this ULIRG sample (Imanishi et al. 2006a). Figures 9d, 9e, and 9f compare the rest-frame equivalent widths of the $3.3 \mu\text{m}$ PAH feature ($\text{EW}_{3.3\text{PAH}}$) with those of the 6.2 , 7.7 , and $11.3 \mu\text{m}$ PAH features.

In Figure 9, the scatter is generally large in plots comparing PAH features at widely separated wavelengths. Contribution from PAH-free AGN continuum to an observed flux can vary to a larger degree with increasing wavelength separation, explaining the trend in Figure 9.

5.3.3. Support for the Strong Dust Temperature Gradient

In a buried AGN with a strong dust temperature gradient, dust extinction toward the 3–4 μm continuum emission regions, estimated using 3–4 μm data, $A_V(3 \mu\text{m})$, should be larger than that toward $10 \mu\text{m}$ continuum emission region estimated from the $\tau'_{9.7}$ value (Fig. 2). For ULIRGs with small contamination from PAH-emitting starbursts to their 3–4 μm spectra (small $\text{EW}_{3.3\text{PAH}}$ values), $A_V(3 \mu\text{m})$ can be estimated with little ambiguity, from absorption optical depths at $3.4 \mu\text{m}$ ($\tau_{3.4}$) by bare carbonaceous dust and at $3.1 \mu\text{m}$ ($\tau_{3.1}$) by ice-covered dust deep inside molecular clouds.

IRAS 08572+3915, 12127–1412, and 17044+6720 show 3–4 μm spectra with very small PAH contaminations. The $\tau_{3.4}$ values are 0.8, 0.35, and 0.15 for IRAS 08572+3915, 12127–1412,

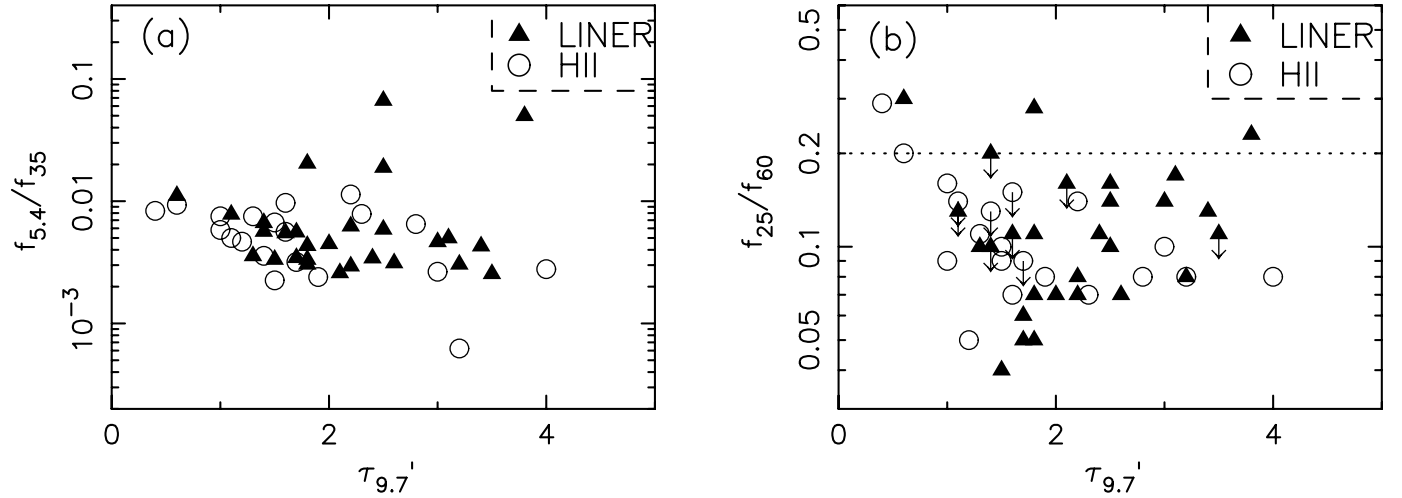


FIG. 13.— Comparison of the strength of the 9.7 μm silicate dust absorption and infrared continuum colors. (a) The abscissa is the $\tau'_{9.7}$ value, and the ordinate is the 5.4–35 μm flux ratio measured from the *Spitzer* IRS spectra in Figs. 3 and 4. Filled triangle: ULIRGs classified optically as LINERs. Open circle: ULIRGs classified optically as H II regions. In the upper right small area separated with the dashed line, there are no data points. (b) The ordinate is the *IRAS* 25–60 μm flux ratio, taken from Table 1. The horizontal dotted line indicates the separation between warm and cool ULIRGs (Sanders et al. 1988b).

and 17044+6720, respectively (Imanishi et al. 2006a). The $\tau_{3.4}$ value reflects the column density of bare carbonaceous dust (without an ice mantle; Pendleton et al. 1994; Imanishi et al. 1996; Rawlings et al. 2003). In the Galaxy, the $\tau'_{9.7}/\tau_{3.4}$ ratio is surprisingly similar in different directions ($\tau'_{9.7}/\tau_{3.4} = 14\text{--}17$; Roche & Aitken 1984, 1985; Pendleton et al. 1994). Thus, a $\tau'_{9.7}/\tau_{3.4}$ ratio substantially smaller than 14–17 can be a signature of a strong dust temperature gradient. Adopting the $\tau'_{9.7}$ values shown in Table 4, we obtain $\tau'_{9.7}/\tau_{3.4} = 5, 7$, and 12 for IRAS 08572+3915, 12127–1412, and 17044+6720, respectively. These three ULIRGs have previously been classified as very strong buried AGN candidates, based on *Spitzer* IRS 5–35 μm spectra (§ 5.2.6). The small $\tau'_{9.7}/\tau_{3.4}$ ratios support the buried AGN hypothesis.

For IRAS 08572+3915 and 17044+6720, no 3.1 μm absorption by ice-covered dust is detected (Imanishi et al. 2006a), suggesting that most of the dust toward the 3–4 μm continuum emission sources is bare, without an ice mantle. However, for IRAS 12127–1412, the 3.1 μm absorption feature is found with $\tau_{3.1} \sim 0.4$, suggesting a large amount of ice-covered dust, in addition to bare dust grains. Since the 9.7 μm absorption is seen whether the silicate dust is bare or ice-covered, adding the column density of ice-covered dust to that of bare carbonaceous dust grains will raise $A_V(3 \mu\text{m})$ even further, strengthening the argument for a strong dust temperature gradient in IRAS 12127–1412.

5.3.4. Classification: AGN or Starburst?

Table 6 (col. [8]) summarizes the detection or nondetection, and the detection significance in the case of detection, of buried AGN signatures based on ground-based 3–4 μm spectra (Imanishi et al. 2006a). The trend of buried AGN detection is generally consistent in individual ULIRGs, in that (1) very strong buried AGN candidates at 3–4 μm are also very strong buried AGN candidates in *Spitzer* IRS 5–35 μm spectra, and (2) ULIRGs with no AGN signatures at 3–4 μm usually show no AGN signatures at 5–35 μm . There are some ULIRGs that have detectable AGN signatures at 5–35 μm , but not at 3–4 μm , which are explained by a larger contamination from modestly obscured starburst emission at a shorter wavelength (Imanishi et al. 2004; Imanishi 2006).

Finally, Imanishi et al. (2006a) have previously argued, based on 3–4 μm spectra, that the fraction of ULIRGs showing lu-

minous buried AGN signatures is higher in LINER ULIRGs than H II region ULIRGs in this sample. This trend is exactly what we have seen in the *Spitzer* IRS 5–35 μm spectra discussed here (§ 5.2.1–5.2.6).

5.4. Buried AGNs and Infrared Colors

Figure 13 compares the strength of the 9.7 μm silicate absorption feature ($\tau'_{9.7}$) and infrared colors, *Spitzer* IRS 5.4–35 μm flux ratio (Fig. 13a), and *IRAS* 25–60 μm flux ratio (Fig. 13b). The infrared colors are independent of $\tau'_{9.7}$ values, with no clear positive or negative correlation found.

The *IRAS* 25–60 μm flux ratio is often used to determine whether ULIRGs show cool ($f_{25}/f_{60} < 0.2$) or warm (> 0.2) far-infrared colors (Sanders et al. 1988b). Although many classical Seyfert-type AGNs surrounded by dusty tori show warm far-infrared colors (de Grijp et al. 1987), the bulk of ULIRGs possessing luminous buried AGN candidates ($\tau'_{9.7} > 2$) belong to cool sources (Fig. 13b). As already argued by Imanishi et al. (2006a) it is quite reasonable that buried AGNs tend to show cooler infrared colors than optical Seyfert AGNs. A larger amount of dust around a central AGN in a buried AGN compared to a Seyfert AGN surrounded by a torus (§ 1) can naturally produce a cooler infrared color, because the contribution from an outer, cooler component becomes important and makes the optical Seyfert signatures fainter or even vanish, because even the torus axis direction can be opaque to the bulk of the AGN's ionizing radiation.

5.5. Future Prospects

Although infrared spectroscopic energy diagnostics enable us to discuss the overall properties of ULIRGs statistically, follow-up observations of infrared-selected buried AGN candidates are very important to cross-check the reliability of these methods. For example, we argue an exceptionally centrally concentrated extreme starburst is unlikely, based on the emission surface brightness and absolute luminosity (§ 5.2.5). X-ray observations can be particularly powerful in distinguishing a buried AGN from an extreme starburst, because a luminous AGN is a much stronger X-ray emitter than any kind of starburst. Since the column densities in the cores of ULIRGs can be large enough to be Compton thick (X-ray absorption $N_H > 10^{24} \text{ cm}^{-2}$), high-sensitivity X-ray observations at $E > 10 \text{ keV}$ are particularly desirable.

In a strongly X-ray-emitting buried AGN surrounded by dense gas and dust, X-ray dissociation regions (XDRs; Maloney et al. 1996) should develop, rather than the photodissociation regions (PDRs) usually seen in a starburst (a strong UV emitter). Finding XDR signatures can thus be another powerful way to test the buried AGN scenario. Imanishi et al. (2004, 2006b) found that several infrared-selected buried AGN candidates tend to show millimeter molecular line flux ratios expected from XDRs. For one of the strongest infrared-selected buried AGN candidates, IRAS 08572+3915, Shirahata et al. (2007) detected CO ($v = 1-0$) absorption features up to $J = 17$ in ground-based infrared 4.5–5 μm observations. It is argued that such high- J lines can be explained only by XDRs (Spaans 2006), further supporting the buried AGN scenario. A further search for XDR signatures in ULIRGs may reinforce the buried AGN picture suggested from the infrared spectroscopy.

6. SUMMARY

We presented *Spitzer* IRS 5–35 μm low-resolution ($R \sim 100$) spectra of ULIRGs at $z < 0.15$ in the *IRAS* 1 Jy sample. ULIRGs classified optically as non-Seyferts (LINERs and H II regions), which comprise 70% of the 1 Jy sample at $z < 0.15$, were our prime targets to search for signatures of optically undetected AGNs deeply buried in dust along virtually all lines of sight, without well-developed narrow-line regions. In total, 28 LINER and 20 H II region ULIRGs (a complete sample) were observed. Three methods were mainly utilized: (1) the equivalent widths of the PAH emission features at 6.2, 7.7, and 11.3 μm , (2) the absolute optical depth of the 9.7 μm silicate dust absorption feature, and (3) the 9.7–18 μm silicate dust absorption optical depth ratio. The first method traces the hardness of the radiation field (starburst-like soft or AGN-like hard). The latter two methods probe the geometry, i.e., whether the dust and energy sources are spatially well mixed (a normal starburst) or the energy source is more centrally concentrated than dust (a buried AGN). We arrived at the following main conclusions.

1. For the majority of ULIRGs, PAH emission features were detected, suggesting the presence of PAH-emitting normal starburst activity. However, the observed PAH-to-infrared luminosity ratios were systematically smaller than in modestly obscured ($A_V < 20$ –30 mag) starburst galaxies. Taken at face value, the PAH-emitting starbursts with modest dust obscuration are energetically insignificant in ULIRGs. The dominant energy sources are (1) very highly obscured ($A_V \gg 30$ mag) PAH-emitting starbursts with mixed dust/source geometry, whose PAH flux is attenuated, and/or (2) non-PAH-emitting AGNs.

2. In 19 out of the 48 (40%) ULIRGs we studied, the observed PAH equivalent widths were found to be substantially lower than normal starburst galaxies at either of the 6.2, 7.7, or 11.3 μm PAH features. This fraction increased to 22/48 (46%) if we adopted the 6.2 μm PAH equivalent widths after correction for the ice and HAC absorption features. The highly obscured energy sources of these ULIRGs' cores were argued to be dominated by non-PAH-emitting AGNs, rather than highly obscured PAH-emitting starbursts.

3. In 20/48 (42%) of ULIRGs, the observed optical depths of the 9.7 μm silicate dust absorption feature ($\tau'_{9.7}$) were substantially larger than the upper limit determined by mixed dust/source geometry in normal starbursts. A non-PAH-emitting buried AGN with centrally concentrated energy source geometry is the

most natural explanation for the energy sources at the cores of these ULIRGs. Comparisons of the $\tau'_{9.7}$ value with the PAH equivalent widths, and with the 11.3–6.2 μm PAH luminosity ratio, supported the buried AGN scenario in these ULIRGs.

4. For selected ULIRGs with small PAH contamination (mostly ULIRGs with large $\tau'_{9.7}$), the 18–9.7 μm silicate absorption optical depth ratios were often significantly smaller than the intrinsic ratio. A strong dust temperature gradient around a centrally concentrated energy source geometry was strongly suggested, further supporting the buried AGN scenario.

5. Sixteen out of 48 (33%) of optical non-Seyfert ULIRGs showed *strong* signatures of luminous buried AGNs, and 26/48 (54%) ULIRGs displayed some signatures. Given that 30% of ULIRGs showed optical Seyfert signatures and were known to possess luminous AGNs surrounded by dusty tori, it was suggested that luminous AGNs are present in at least 50%, and possibly >65%, of the complete ULIRG sample at $z < 0.15$.

6. The fraction of ULIRGs showing luminous buried AGN signatures was always higher in optical LINER ULIRGs than H II region ULIRGs using any of the above methods. This was qualitatively explained by the scenario that the starburst regions probed in the optical are dustier in LINER ULIRGs than in H II region ULIRGs.

7. For ULIRGs with strong buried AGN signatures and small PAH contamination, the absorption-corrected intrinsic luminosities of the buried AGNs were found to be quantitatively comparable to $10^{12} L_\odot$, accounting for the bulk of the observed infrared luminosities of ULIRGs.

8. Most of the luminous buried AGN candidates were found in ULIRGs with cool far-infrared colors.

9. The above conclusions drawn from the *Spitzer* IRS 5–35 μm spectra were generally consistent with those based on 3–4 μm spectra obtained by Imanishi et al. (2006a). Our main conclusions were that optically undetected luminous buried AGNs are surely present in a significant fraction of non-Seyfert ULIRGs. It is therefore very important to carefully and quantitatively investigate the energetic roles of optically elusive luminous buried AGNs, if we are to understand the true nature of the nearby ULIRG population.

We are grateful to J. E. Chiar for providing M. I. with silicate dust absorption profiles of the Galactic sources in an electronic format, and H. Suto for valuable discussions about the properties of silicate dust grains. We thank V. Charmandaris for the advice about the *Spitzer* IRS data analysis and the anonymous referee for valuable comments. M. I. is supported by Grants-in-Aid for Scientific Research (16740117). This work is based on observations made with the *Spitzer Space Telescope*, which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under a contract with NASA. Support for this work was provided by NASA and also by an award issued by JPL/Caltech. Research in Infrared Astronomy at the Naval Research Laboratory is supported by the Office of Naval Research (USA). This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France, and of the NASA/IPAC Extragalactic Database (NED), which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under contract with the National Aeronautics and Space Administration.

APPENDIX

THE MAXIMUM $\tau'_{9.7}$ VALUE IN A MIXED DUST/SOURCE GEOMETRY

Figure 1a shows a schematic diagram of a typical mixed dust/source geometry. In this geometry, the observed flux $I(\tau_\nu)$ is given by

$$I(\tau_\lambda) = I_0(\lambda) \times \frac{1 - e^{-\tau_\lambda}}{\tau_\lambda}, \quad (\text{A1})$$

where $I_0(\lambda)$ is an unattenuated intrinsic flux and τ_λ is the optical depth at each wavelength, which takes different values inside absorption features and nearby continua outside the features.

We here consider radiation with an intrinsically smooth spectral energy distribution, namely $I_0(8 \mu\text{m}) \sim I_0(9.7 \mu\text{m}) \sim I_0(13 \mu\text{m})$, such as expected for dust emission in ULIRGs. The peak absorption optical depth of the $9.7 \mu\text{m}$ silicate feature, relative to the nearby continuum at 8 and $13 \mu\text{m}$ (which corresponds to $\tau'_{9.7}$ and not $\tau_{9.7}$ in our definition in § 4.2), is

$$\tau'_{9.7} \equiv \ln \left(\frac{1 - e^{-\tau_{\text{cont}}}}{\tau_{\text{cont}}} \frac{\tau_{97\text{sil}}}{1 - e^{-\tau_{97\text{sil}}}} \right) \quad (\text{A2})$$

$$= \ln \left(\frac{1 - e^{-\tau_{\text{cont}}}}{1 - e^{-\tau_{97\text{sil}}}} \frac{\tau_{97\text{sil}}}{\tau_{\text{cont}}} \right), \quad (\text{A3})$$

where $\tau_{97\text{sil}}$ and τ_{cont} refer to the optical depths at the peak of the $9.7 \mu\text{m}$ silicate absorption feature and nearby continuum at 8 and $13 \mu\text{m}$ just outside the feature, respectively. The $\tau_{97\text{sil}}$ value is always larger than the τ_{cont} value, because $\tau_{97\text{sil}}$ is the superposition of the optical depth for continuum emission and the $9.7 \mu\text{m}$ silicate feature.

In the case of the Galactic dust extinction curve derived by Rieke & Lebofsky (1985),

$$\tau_{\text{cont}} = 0.025/1.08A_V, \quad (\text{A4})$$

$$\tau_{97\text{sil}} = 0.087/1.08A_V, \quad (\text{A5})$$

where, by definition, $A_V \equiv 1.08 \times \tau_V$.

The absorption strength of the $9.7 \mu\text{m}$ silicate feature, relative to A_V , is different by a factor of ~ 2 in different directions in the Galaxy (Roche & Aitken 1984, 1985). If we take this into account,

$$\tau_{\text{cont}} = 0.025/1.08A_V, \quad (\text{A6})$$

$$\tau_{97\text{sil}} = (0.05-0.1)A_V + 0.025/1.08A_V. \quad (\text{A7})$$

The τ_{cont} and $\tau_{97\text{sil}}$ values can be arbitrarily large with increasing amount of dust (A_V) in the mixed dust/source geometry. The $\tau'_{9.7}$ value is an increasing function of A_V , and the first term in the equation (A3) $[(1 - e^{-\tau_{\text{cont}}})/(1 - e^{-\tau_{97\text{sil}}})]$ becomes unity for a very large A_V value. Thus, the upper limit to $\tau'_{9.7}$ is

$$\tau'_{9.7}(\text{upper limit}) = \ln \left(\frac{\tau_{97\text{sil}}}{\tau_{\text{cont}}} \right). \quad (\text{A8})$$

We obtain a stringent upper limit of $\tau'_{9.7} \sim 1.7$ in this mixed dust/source geometry for any amount of dust.

The τ_{cont}/A_V and $\tau_{97\text{sil}}/A_V$ ratios can vary substantially with varying dust grain size, because the wavelength difference between the optical V band and infrared $\sim 10 \mu\text{m}$ is large. However, the upper limit of $\tau'_{9.7}$ in the mixed dust/source geometry is dependent only on $\tau_{\text{cont}}/\tau_{97\text{sil}}$, which should be less sensitive to varying dust grain size (see Fig. 6b), because their wavelengths are very close. Therefore, the above conclusion about the upper limit of $\tau'_{9.7}$ in the mixed dust/source geometry should be robust to possible differences in dust grain size between the cores of ULIRGs and the Galactic diffuse interstellar medium.

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